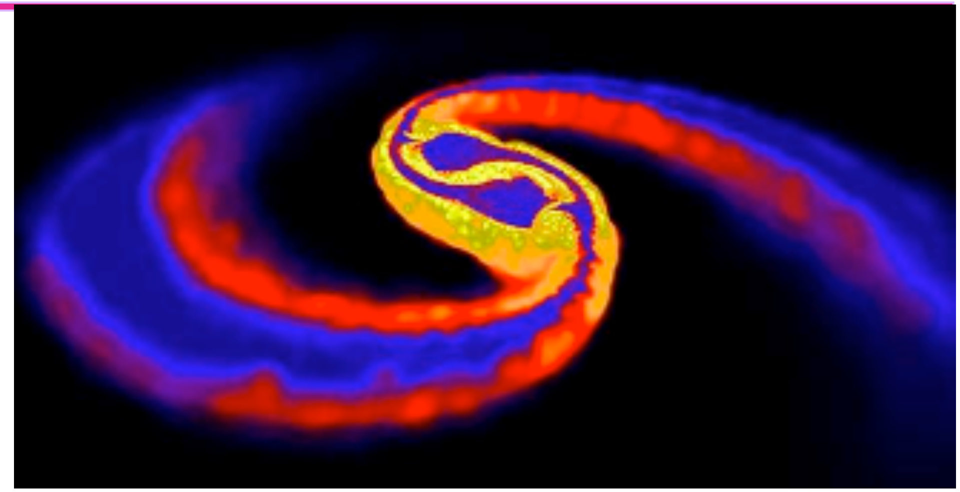
First Results from Advanced LIGO's Search for Gravitational Waves

- Gravitational waves
- Survey of astrophysical sources and signal morphologies
- Compact Binary Coalescences
- What we can learn: physics and astrophysics



"Merging Neutron Stars" (Price & Rosswog)

Oh wait … yes, there is!

Alan Weinstein, Caltech

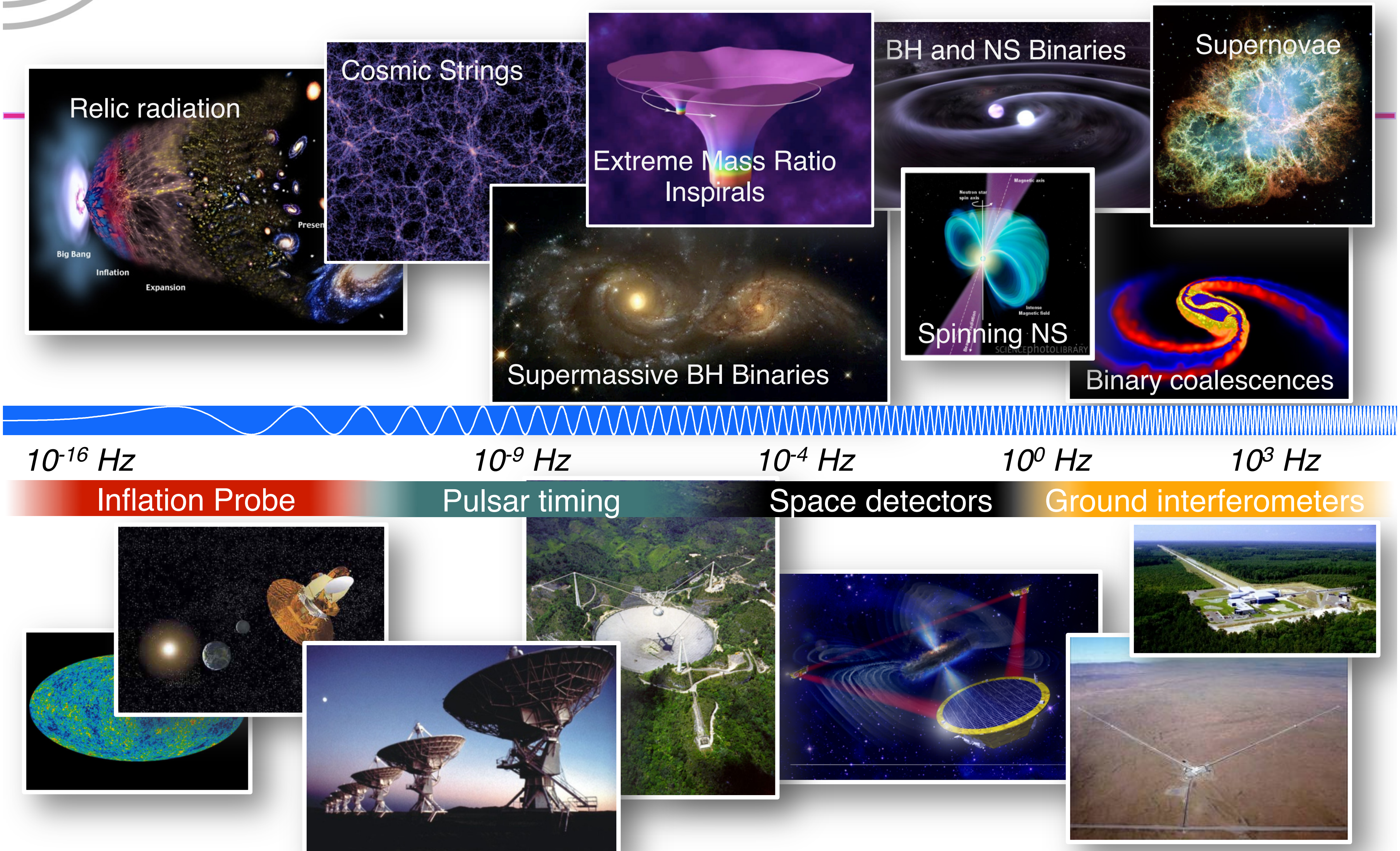
for the LIGO Scientific Collaboration

LIGO-G1600287

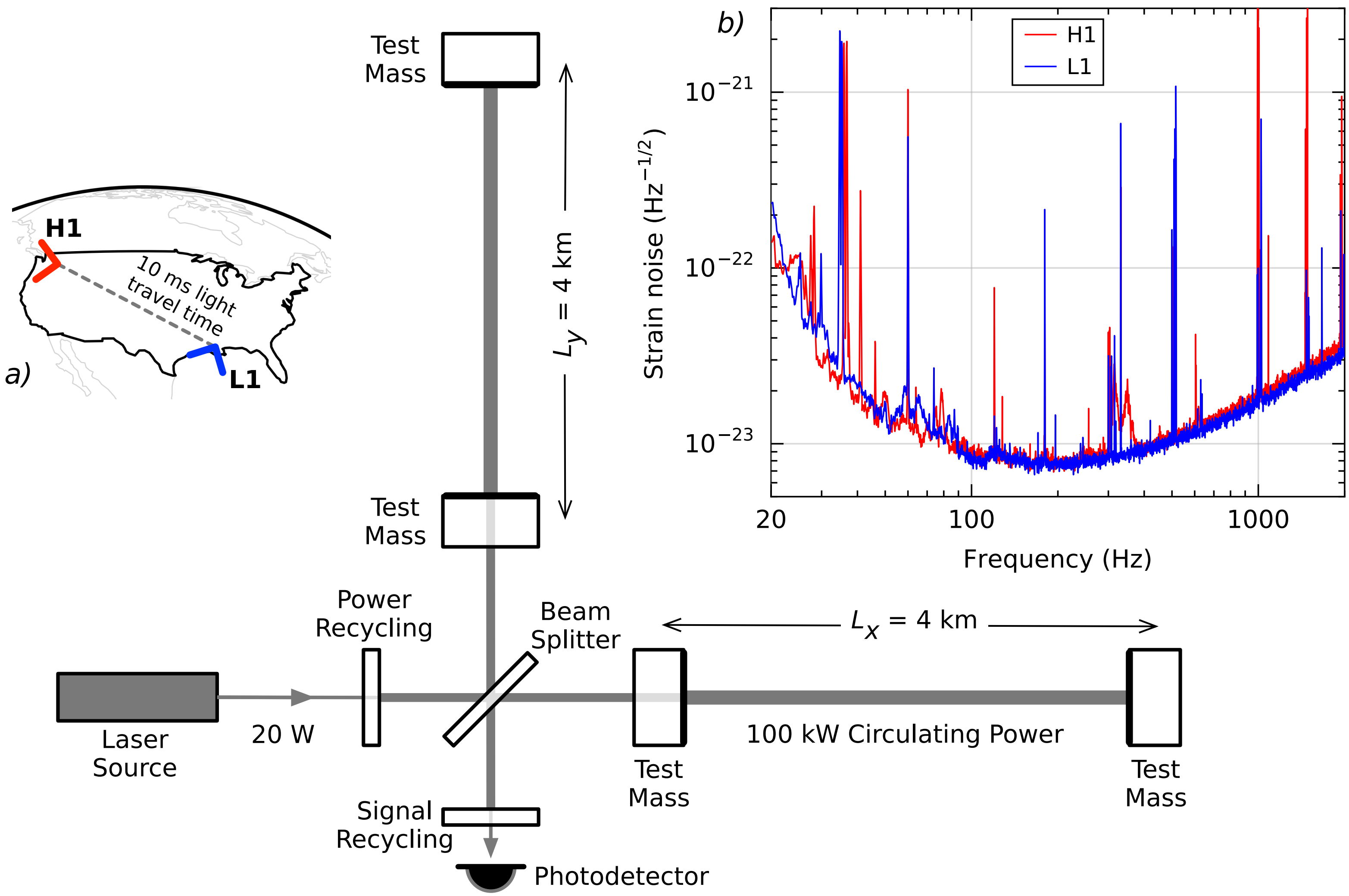
Fig. 1.1 – LIGO detector with 4 km arms at Livingston, Louisiana

Fig. 1.2 – Virgo Detector, with 3 km arms, at Cascina, near Pisa

The GW Spectrum

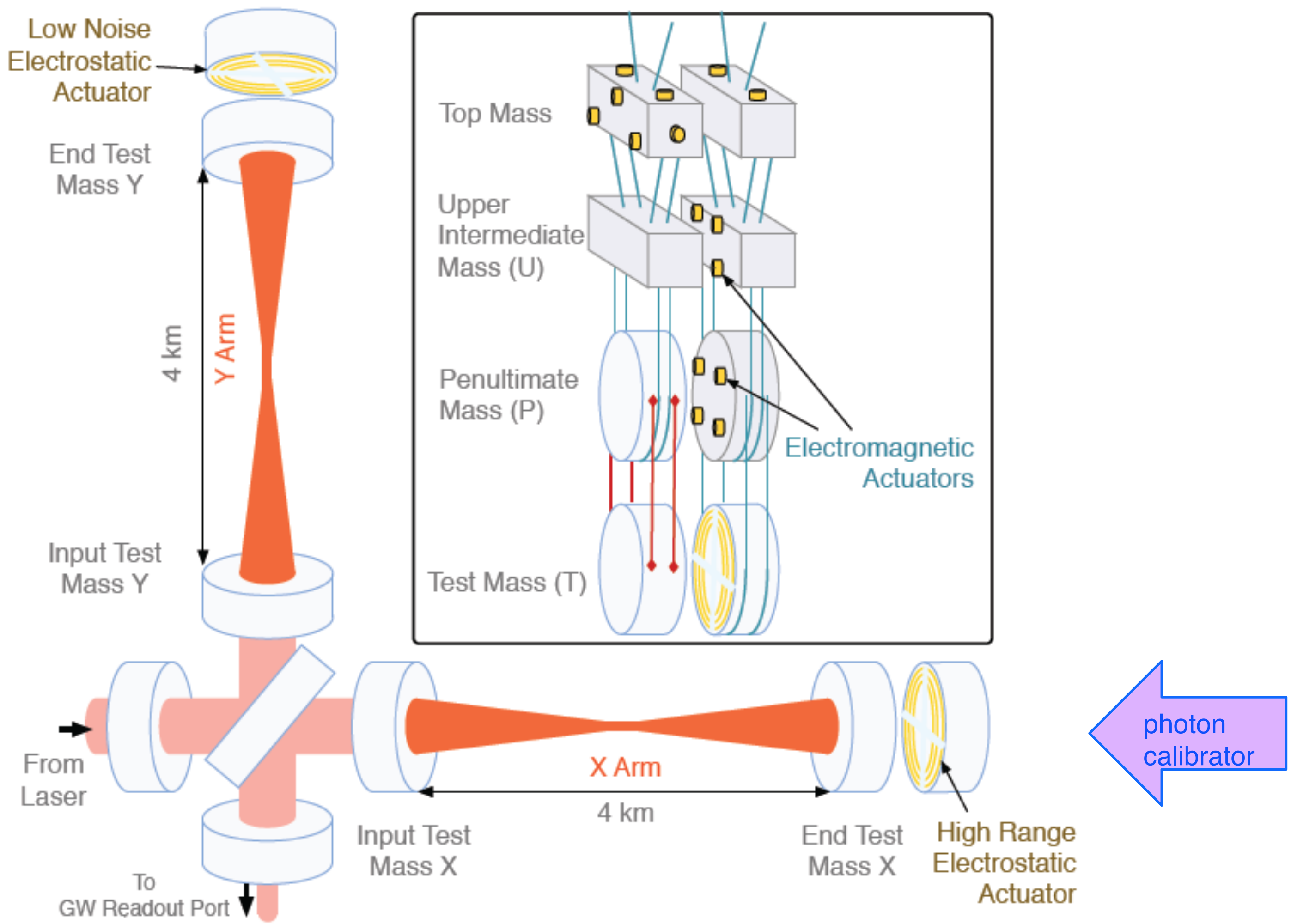


The Advanced LIGO detectors



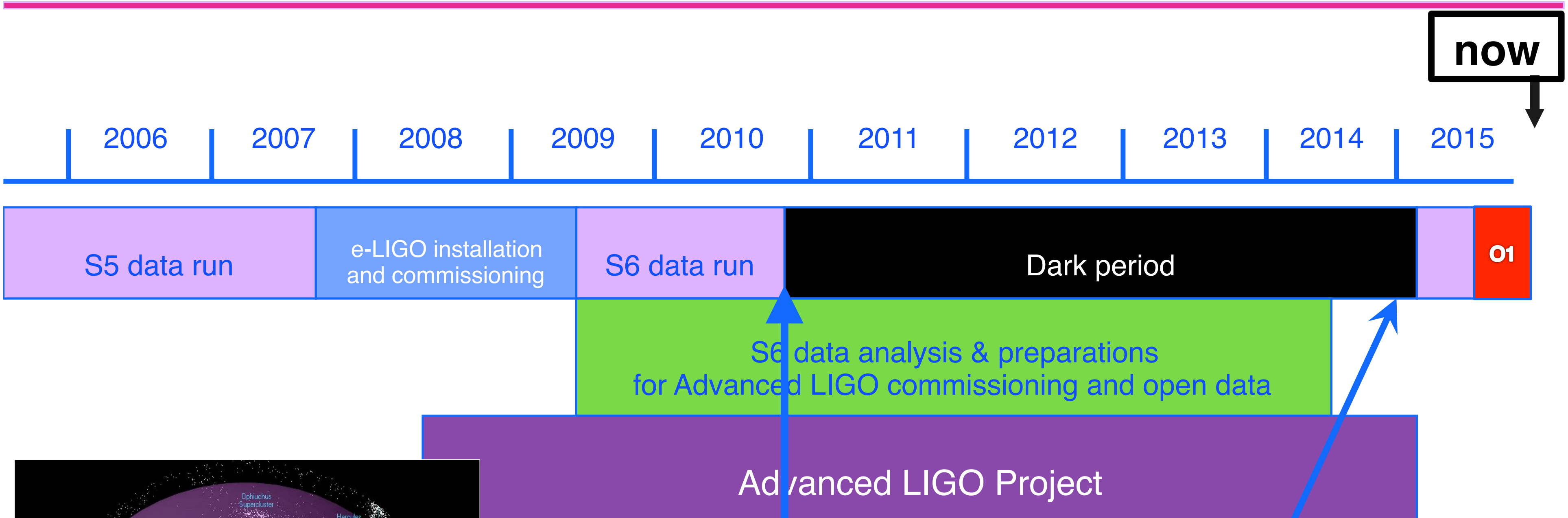
https://dcc.ligo.org/LIGO-P1500237/public/main

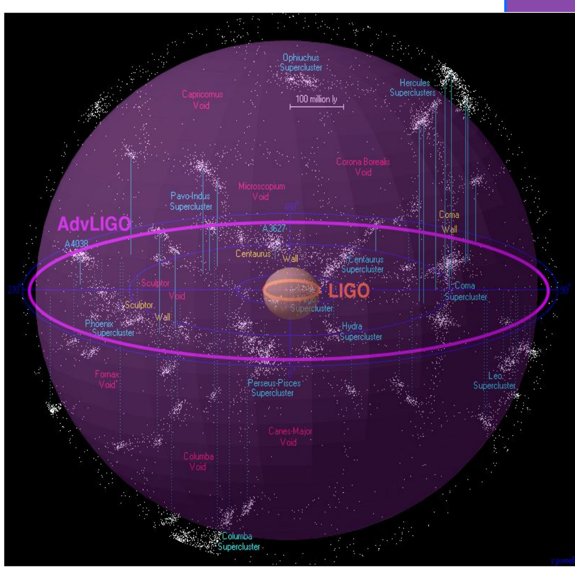
End mirror (“test mass”) quadruple-pendulum suspensions



https://dcc.ligo.org/LIGO-P1500248/public/main

iLIGO → eLIGO → aLIGO





Improve amplitude sensitivity by a factor of 10x, and…

⇒ Number of sources goes up 1000x!

LIGO Scientific Collaboration

Abilene Christian University
Albert-Einstein Institut
Andrews University
American University
California Institute of Technology
California State Univ., Fullerton
Canadian Inst. Th. Astrophysics
Carleton College
College of William and Mary
Columbia U. in the City of New York
Embry-Riddle Aeronautical Univ.
Eötvös Loránd University
Georgia Institute of Technology
Goddard Space Flight Center
Hobart & William Smith Colleges
ICTP-SAIFR
IndIGO
IAP-Russian Acad. of Sciences
Inst. Nacional Pesquisas Espaciais
Kenyon College
Korean Gravitational-Wave Group
Louisiana State University
Montana State University
Montclair State University
Moscow State University
National Tsinghua University
Northwestern University

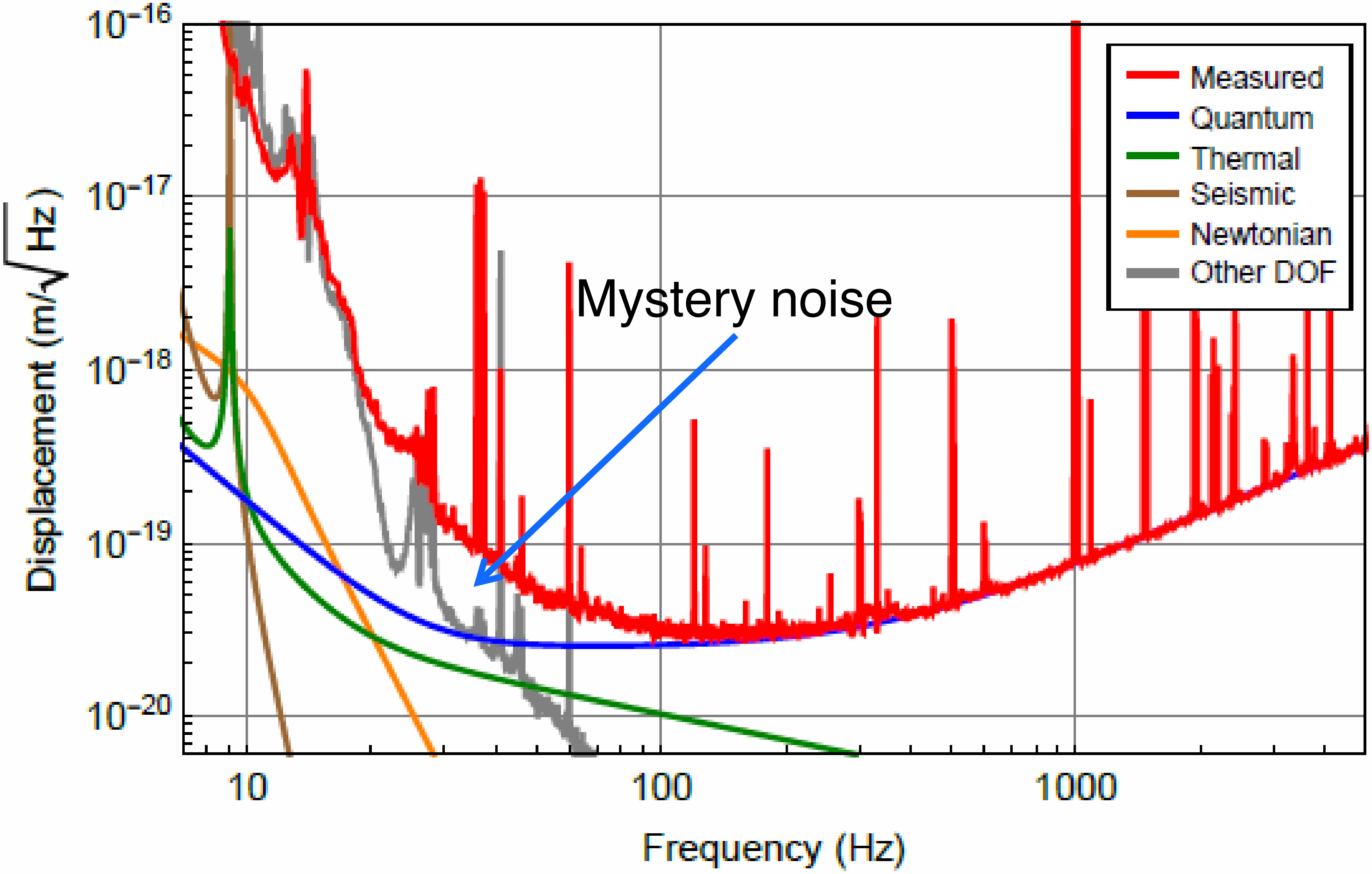
Penn State University
Rochester Institute of Technology
Sonoma State University
Southern Univ. and A&M College
Stanford University
Syracuse University
Szeged University
Texas Tech University
Trinity University
Tsinghua University
Universitat de les Illes Balears
University of Alabama in Huntsville
University of Brussels
University of Chicago
University of Florida
University of Maryland
University of Michigan
University of Minnesota
University of Mississippi
University of Oregon
University of Sannio
Univ. of Texas-Rio Grande Valley
University of Washington
University of Wisconsin-Milwaukee
Washington State University
West Virginia University
Whitman College

LIGO Laboratory: California Institute of Technology, Massachusetts Institute of Technology, LIGO Hanford Observatory, LIGO Livingston Observatory

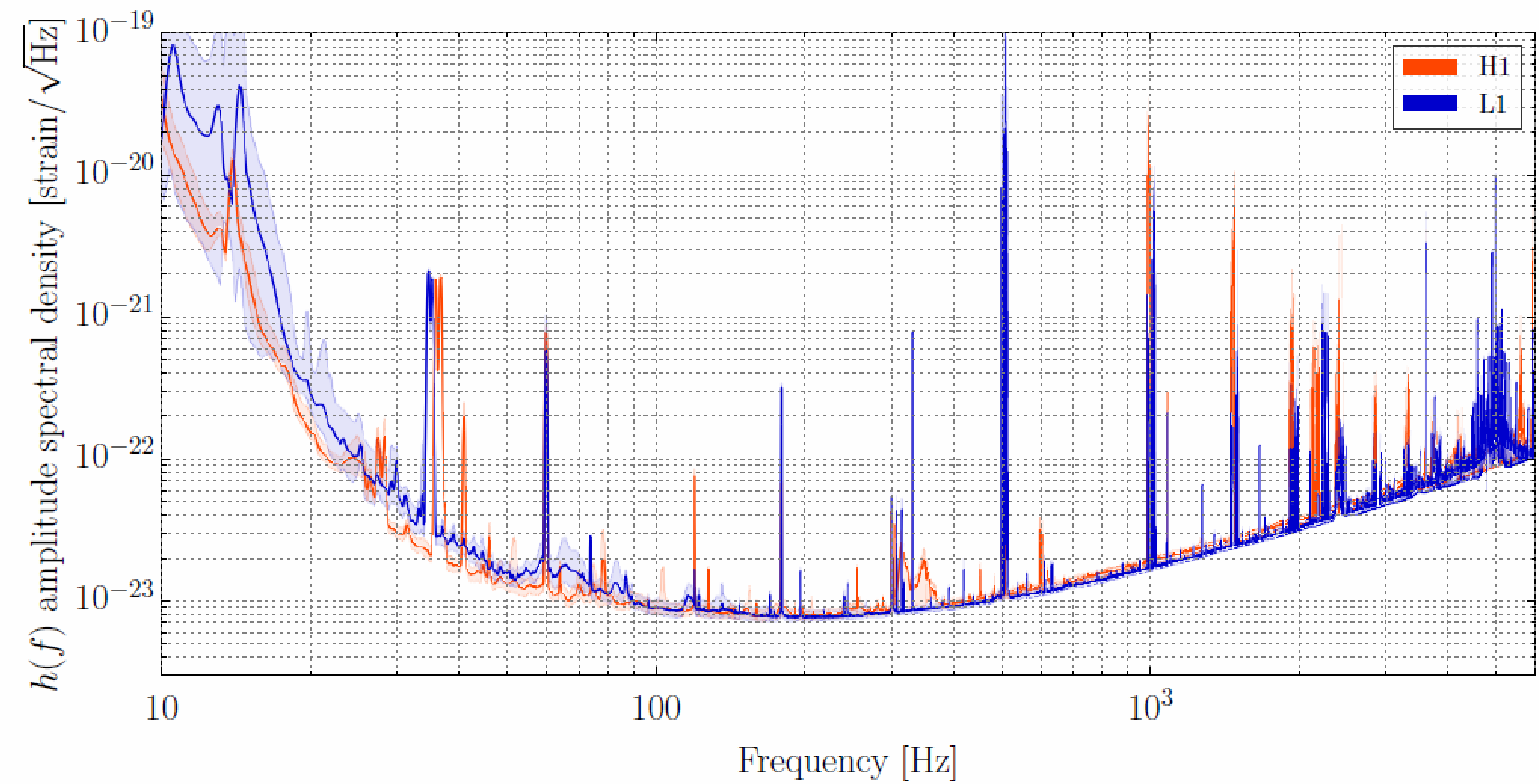
Australian Consortium for Interferometric Gravitational Astronomy (ACIGA):
Australian National University, Charles Sturt University, Monash University, University of Adelaide, University of Melbourne, University of Western Australia

German/British Collaboration for the Detection of Gravitational Waves (GEO600):
Cardiff University, Leibniz Universität Hannover, Albert-Einstein Institut, Hannover, King's College London, Rutherford Appleton Laboratory, University of Birmingham, University of Cambridge, University of Glasgow, University of Hamburg, University of Sheffield, University of Southampton, University of Strathclyde, University of the West of Scotland

Noise sources that limit performance during O1

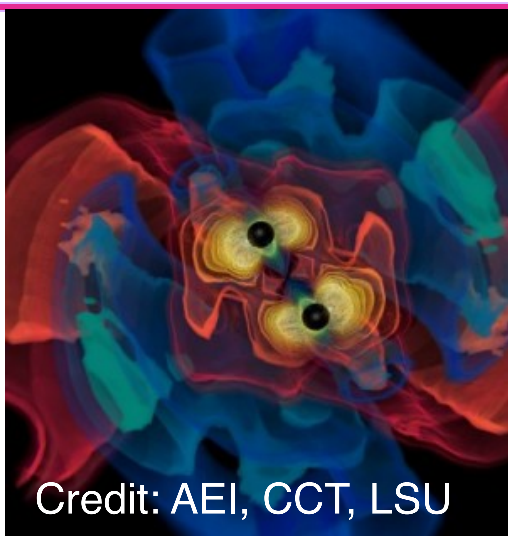


Non-stationarity of the noise ASD



https://dcc.ligo.org/LIGO-P1500238/public/main

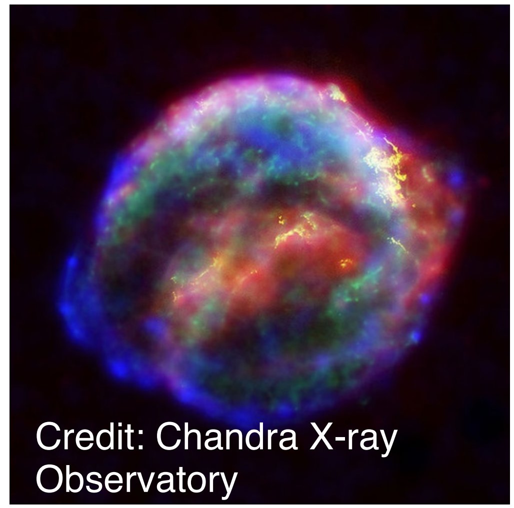
GW sources for ground-based detectors: The most energetic processes in the universe


Credit: AEI, CCT, LSU

<u>Coalescing Compact Binary Systems</u>: Neutron Star-NS, Black Hole-NS, BH-BH

- Strong emitters, well-modeled,

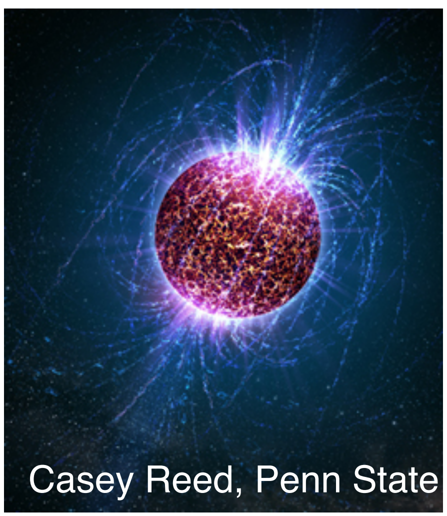
- (effectively) transient


Credit: Chandra X-ray Observatory

<u>Asymmetric Core Collapse Supernovae</u>

- Weak emitters, not well-modeled ('bursts'), transient

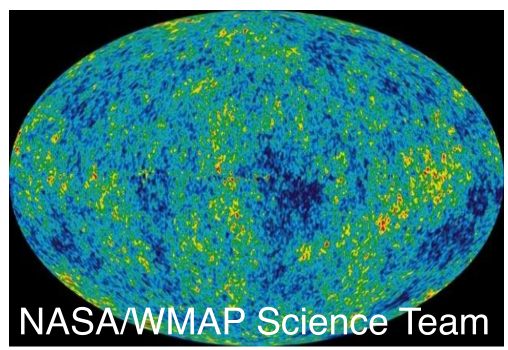
- Cosmic strings, soft gamma repeaters, pulsar glitches also in 'burst' class


Casey Reed, Penn State

<u>Spinning neutron stars</u>

- (effectively) monotonic waveform

- Long duration

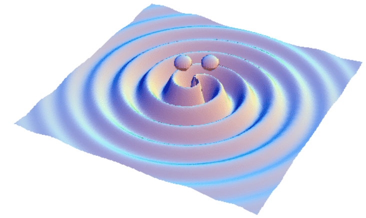

NASA/WMAP Science Team

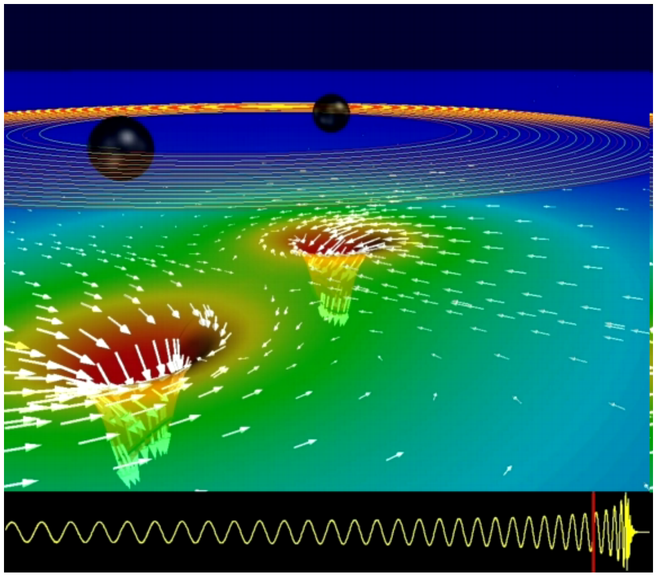
<u>Cosmic Gravitational-wave Background</u>

- Residue of the Big Bang, long duration

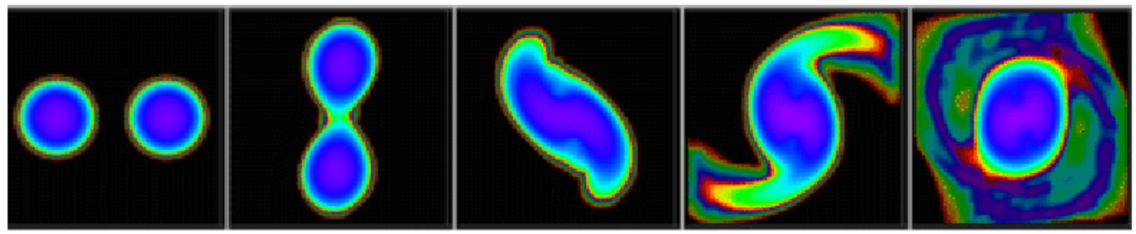
- Long duration, stochastic background

GWs from coalescing compact binaries (NS/NS, BH/BH, NS/BH)





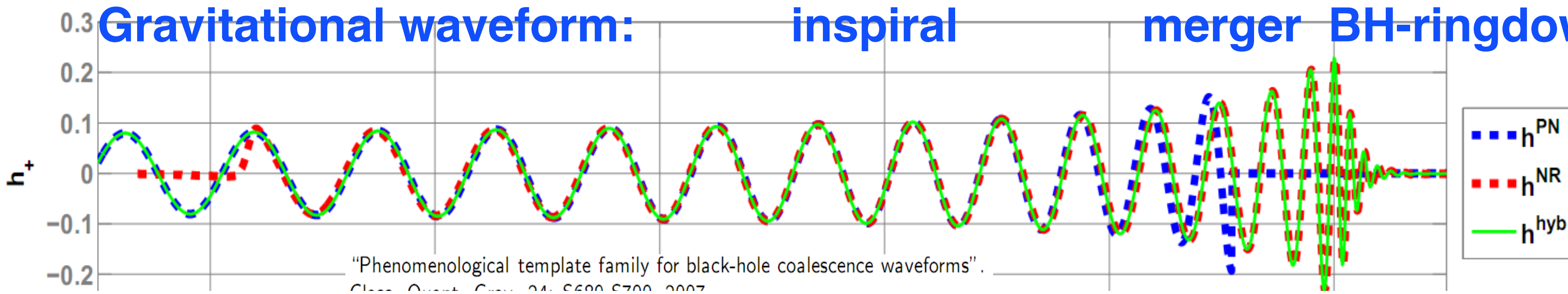
- Neutron star – neutron star (Centrella et al.)



Tidal disruption of neutron star

A unique and powerful laboratory to study strong-field, highly dynamical gravity and the structure of nuclear matter in the most extreme conditions

Gravitational waveform: inspiral merger BH-ringdown

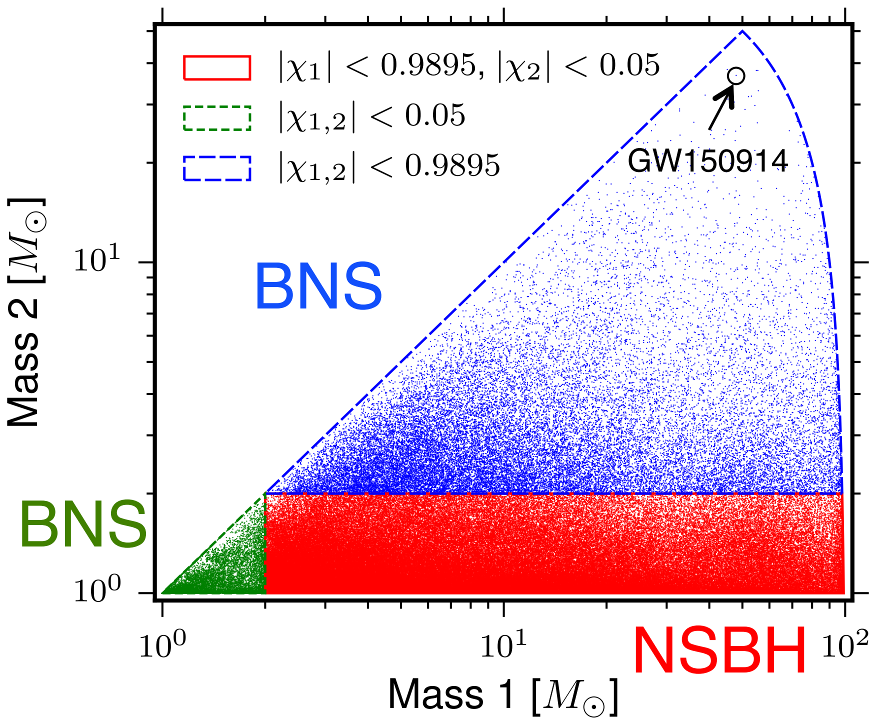


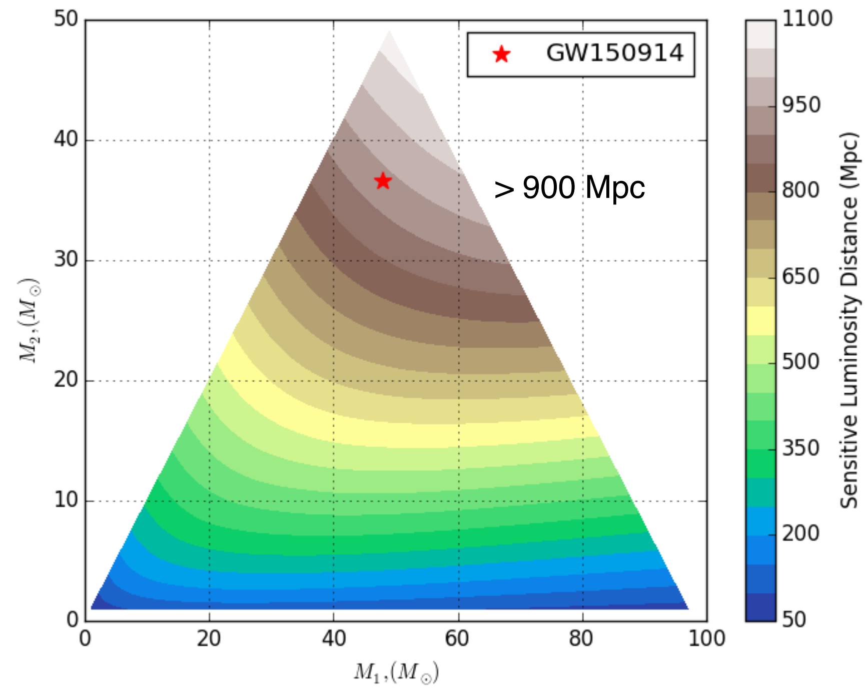
Waveform carries lots of information about binary masses, orbit, merger

Search pipelines

- We need to sift through months of two-detector-coincident strain data h(t) to look for signals with durations from minutes to fractions of a second, above the detector noise.
- **Two different template-based searches for compact binary coalescence (CBC): BNS, NSBH,BBH**:
  - Low-latency (10's of seconds) – gstlal (gstreamer-based)
  - "Offline" - pyCBC (fft-based)
- **Two different searches for short-duration, unmodeled "bursts" of GW power in the time-frequency plane, with low latency:**
  - Coherent WaveBurst - cWB
  - Online LIGO Inference Burst – oLIB
- All make use of **two-detector coincidence in time and in signal morphology**.
- All estimate the background from accidental coincidence of instrumental noise triggers, using "time slides" or variations thereof.
- **All detected GW150914 with high significance above detector noise**

Template-based searches





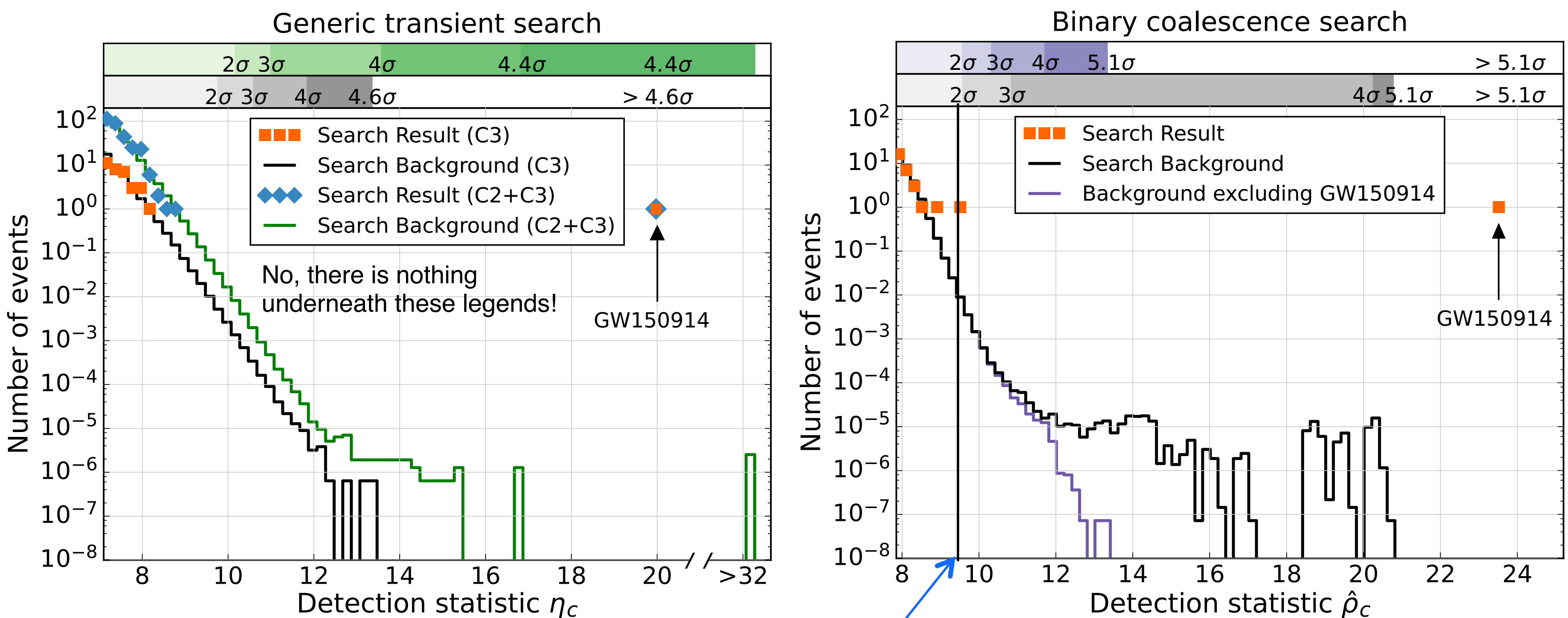
Masses and (aligned) spins
Templates spaced for < 3% loss of SNR: 250K templates.

Sensitive distance in Mpc

https://dcc.ligo.org/LIGO-P1500269/public/main

Results of search over first 38 days of Observation run 1

(Sep 12 – Oct 20, 2015)



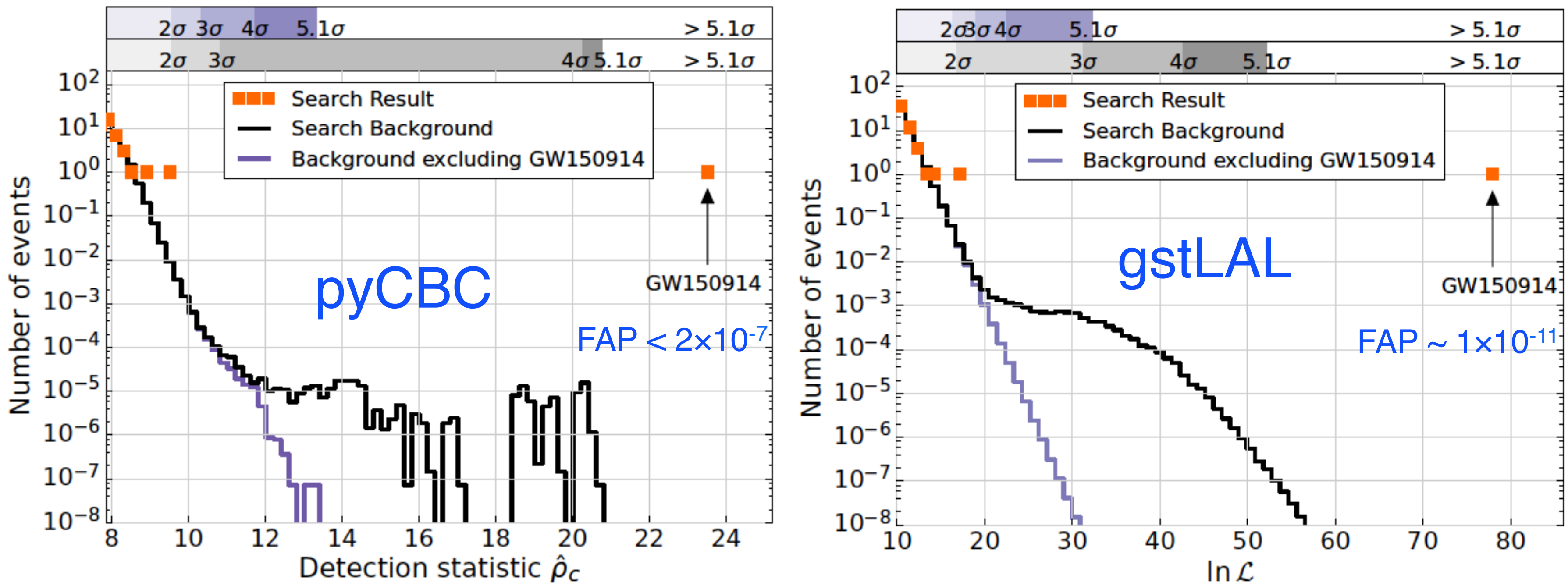
Is it likely that the first detected event should be so loud?

P(1 event with $\rho \geq 23.7$) = $(9.5/23.7)^3$ = 6%.

Are there fainter events? Yes, one! LVT151012, $\rho = 9.6$.

https://dcc.ligo.org/LIGO-P1500269/public/main

Both CBC pipelines detect signal with high significance



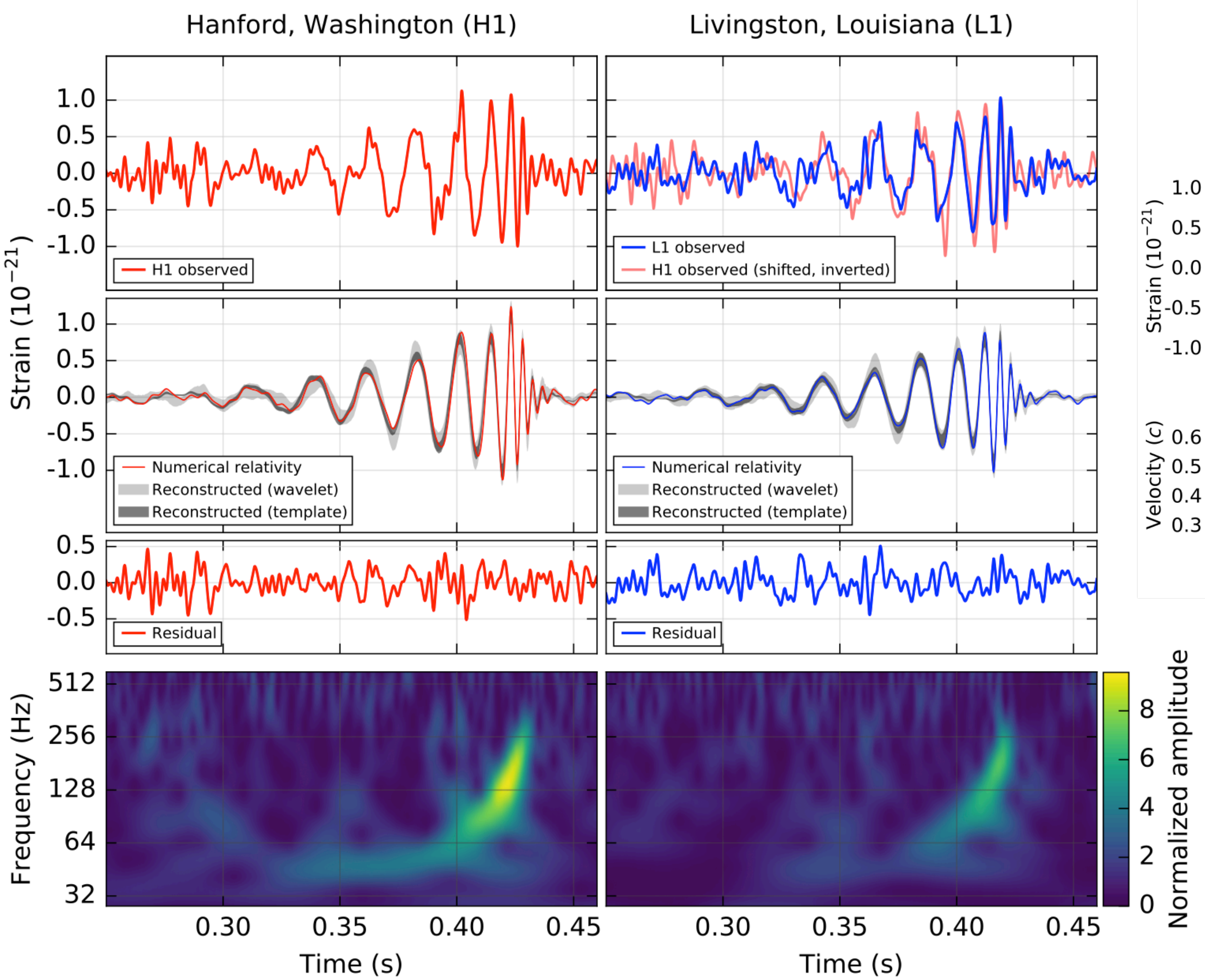
Caltech LIGO: Surabhi Sachdev, Kent Blackburn, AJW
Visitors: Duncan Meacher (PSU), Cody Messick (PSU)

LVT151012	12 October 2015 09:54:43	0.44	0.02 (2.1 σ)	15^{+1}_{-1}	23^{+18}_{-5}	13^{+4}_{-5}	$0.0^{+0.3}_{-0.2}$	1100^{+500}_{-500}

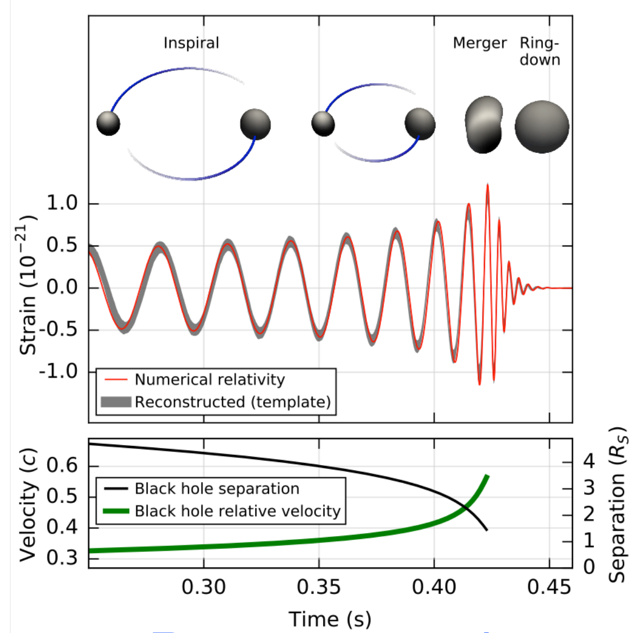
https://dcc.ligo.org/LIGO-P1500269/public/main

GW150914

Phys. Rev. Lett. 116, 061102 – Published 11 February 2016
https://dcc.ligo.org/LIGO-P150914/public/main



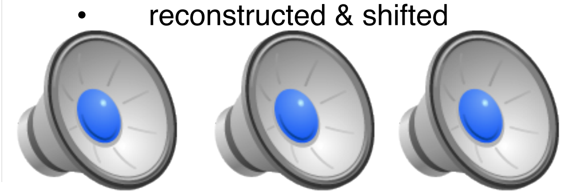
Whitened and band-passed [40-300] Hz



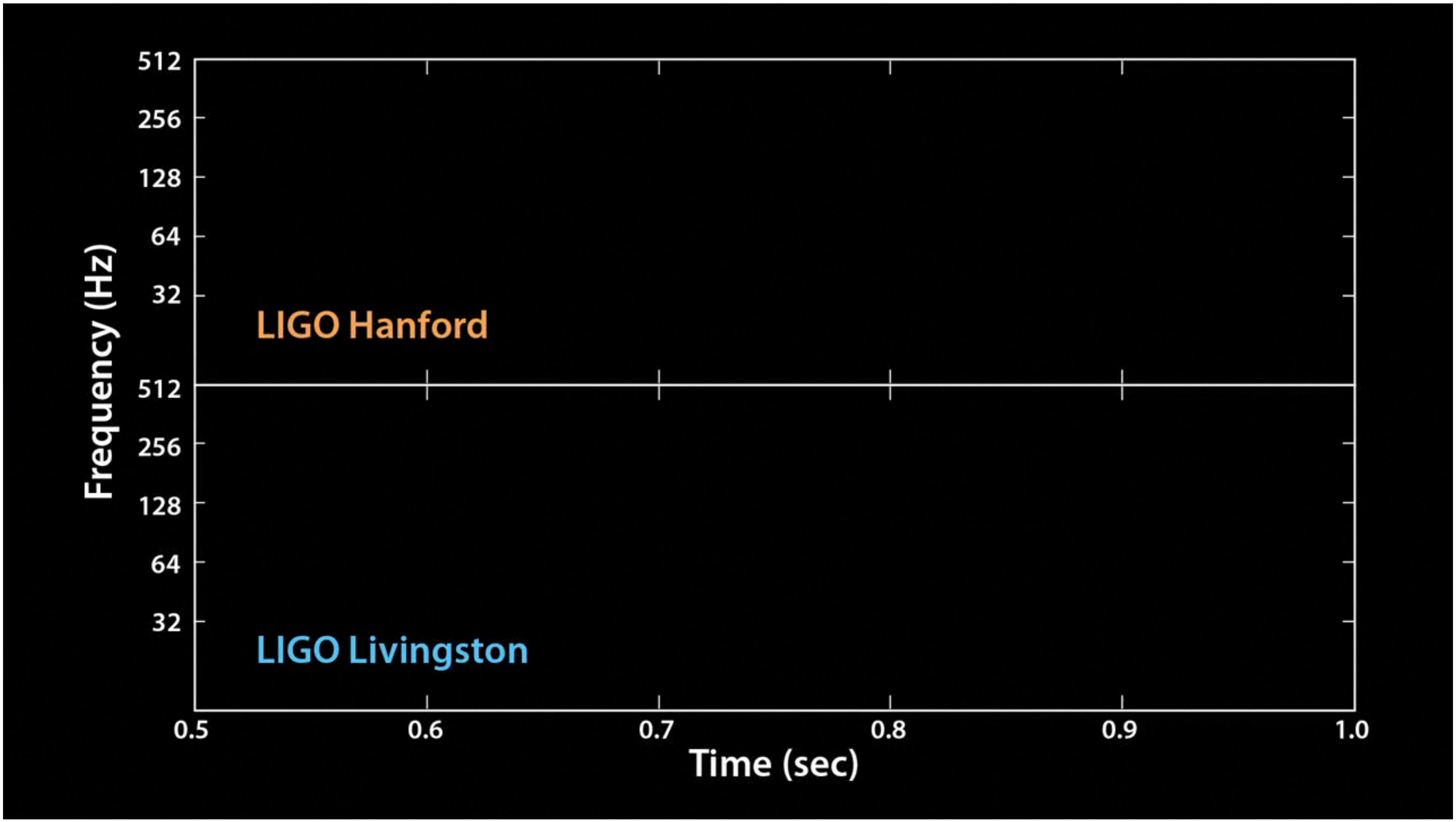
Reconstructed (no whitening)

Audio:

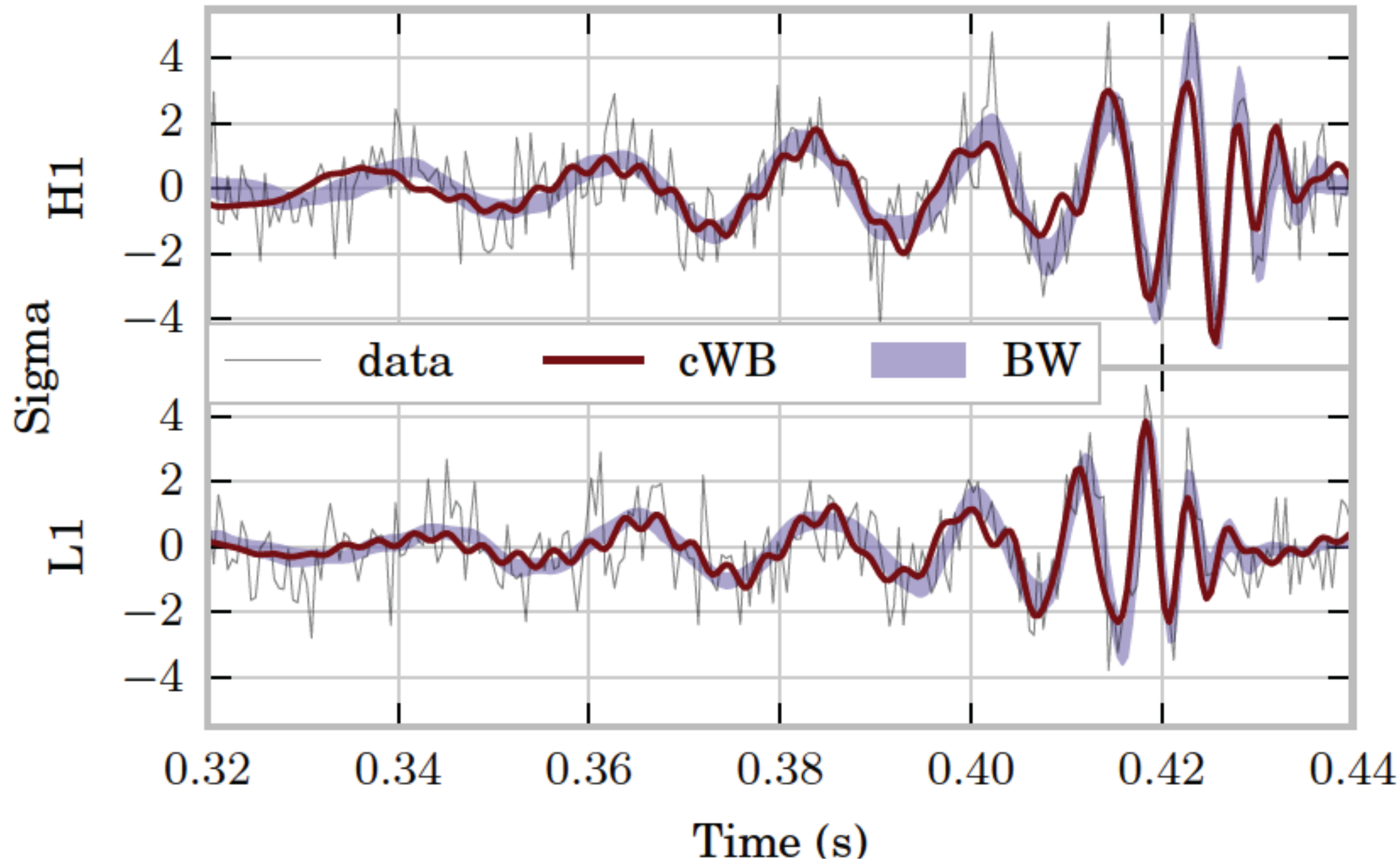
- filtered data
- freq-shifted data
- reconstructed & shifted



The sound of two black holes merging



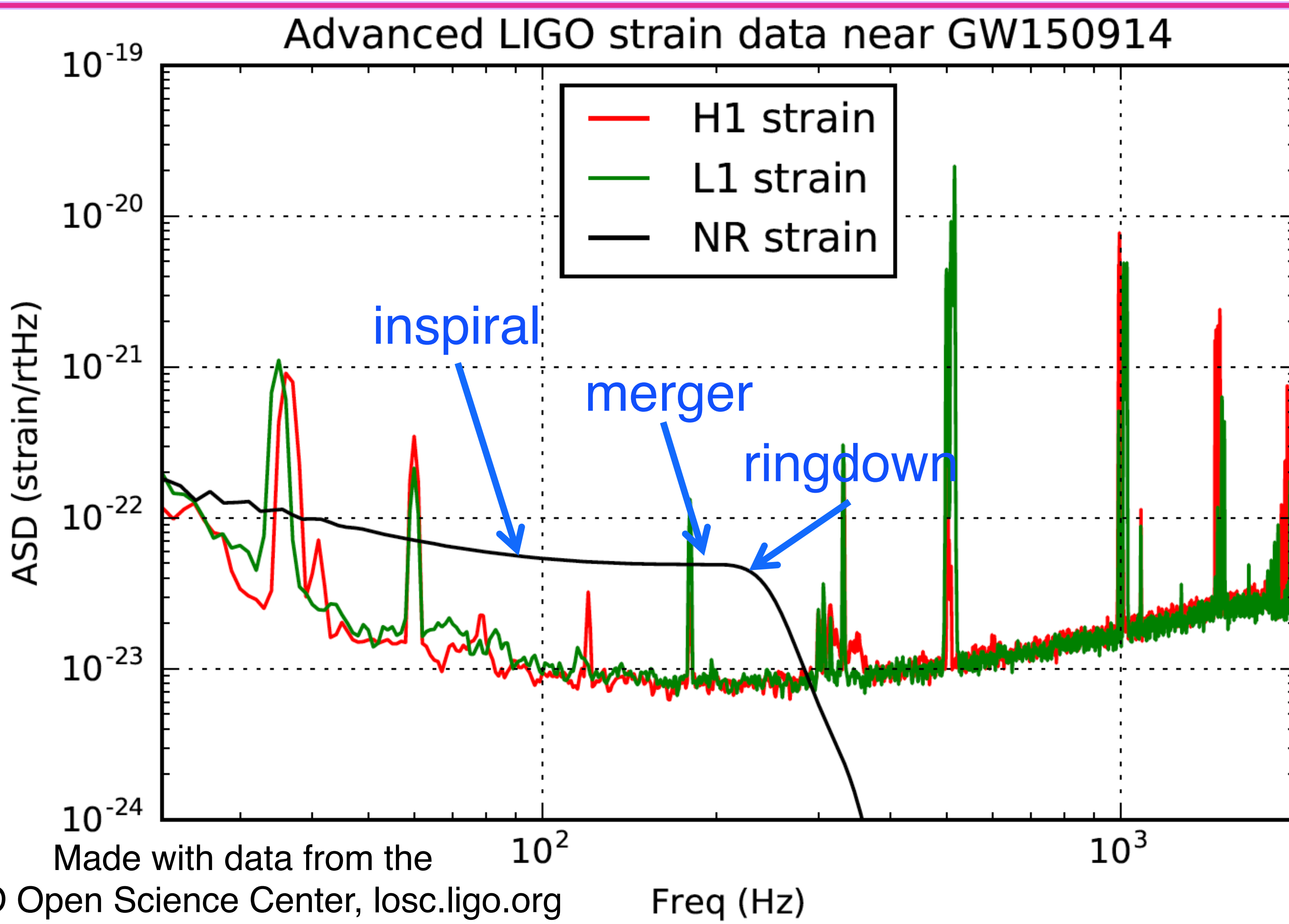
Unmodeled ("burst") signal reconstruction



Caltech LIGO: Jonah Kanner

https://dcc.ligo.org/LIGO-P1500229/public/main

GW150914 in the frequency domain



Made with data from the LIGO Open Science Center, losc.ligo.org

The LIGO Open Science Center
losc.ligo.org

LIGO Open Science Center

LIGO is operated by California Institute of Technology and Massachusetts Institute of Technology and supported by the U.S. National Science Foundation.

- Getting Started
- Tutorials
- Data & Catalogs
- Timelines
- My Sources
- Software
- GPS ↔ UTC
- About LIGO
- Student Projects
- Acknowledgement

Data release for event GW150914

This page has been prepared by the LIGO Scientific Collaboration (LSC) and the Virgo Collaboration to inform the broader community about a confirmed astrophysical event observed by the gravitational-wave detectors, and to make the data around that time available for others to analyze. There is also a **technical details** page about the data linked below, and feel free to **contact us**. This dataset has the Digital Object Identifier (doi) http://dx.doi.org/10.7935/K5MW2F23

Summary of Observation

The event occurred at GPS time 1126259462.39 == September 14 2015, 09:50:45.39 UTC. The false alarm rate is estimated to be less than 1 event per **203,000 years**, equivalent to a significance of **5.1 sigma**. The event was detected in data from the LIGO Hanford and LIGO Livingston observatories.

- There are Science Summaries, covering the information below in ordinary language.
- There is a one page factsheet about GW150914, summarizing the event.

How to Use this Page

- **Click on the section headings below to show available data files.**
  - (click to Open/Close all sections)
- There are lots of data files available in the sections below, look for the word **DATA**.
- Click on each thumbnail image for larger image.
- See the papers linked below for full information, references, and meaning.
- Many of the data files linked below have heterogeneous formatting; if you have any questions, please contact us.

The G150914 detection paper:

Observation of Gravitational Waves from a Binary Black Hole Merger

The data from the observatories from which the science is derived:

Gravitational-Wave Strain Data

- **Tutorial on Signal Processing with Gravitational-Wave Strain Data**
- **About the Instruments and Collaborations**
- **Observing Gravitational-Wave Transient GW150914 with Minimal Assumptions**
- **GW150914: First Results from the Search for Binary Black Hole Coalescence with Advanced LIGO**
- **Properties of the binary black hole merger GW150914**
- **The Rate of Binary Black Hole Mergers Inferred from Advanced LIGO Observations Surrounding GW150914**
- **Astrophysical Implications of the Binary Black-Hole Merger GW150914**
- **Tests of general relativity with GW150914**
- **GW150914: Implications for the Stochastic Gravitational-Wave Background from Binary Black Holes**
- **High-energy Neutrino Follow-up Search of Gravitational Wave Event GW150914 with IceCube and ANTARES**
- **GW150914: The Advanced LIGO Detectors in the Era of First Discoveries**

SIGNAL PROCESSING WITH GW150914 OPEN DATA

Welcome! This ipython notebook (or associated python script GW150914_tutorial.py) will go through some typical signal processing tasks on strain time-series data associated with the LIGO GW150914 data release from the LIGO Open Science Center (LOSC):

- https://losc.ligo.org/events/GW150914/
- View the tutorial as a web page - https://losc.ligo.org/s/events/GW150914/GW150914_tutorial.html
- Download the tutorial as a python script - https://losc.ligo.org/s/events/GW150914/GW150914_tutorial.py
- Download the tutorial as iPython Notebook - https://losc.ligo.org/s/events/GW150914/GW150914_tutorial.ipynb

To begin, download the ipython notebook, readligo.py, and the data files listed below, into a directory / folder, then run it. Or you can run the python script GW150914_tutorial.py. You will need the python packages: numpy, scipy, matplotlib, h5py.

On Windows, or if you prefer, you can use a python development environment such as Anaconda (https://www.continuum.io/why-anaconda) or Enthought Canopy (https://www.enthought.com/products/canopy/).

Questions, comments, suggestions, corrections, etc: email losc@ligo.org

v20160208b

Intro to signal processing

This tutorial assumes that you know python well enough.

If you know how to use "ipython notebook", use the GW150914_tutorial.ipynb file. Else, you can use the GW150914_tutorial.py script.

This tutorial assumes that you know a bit about signal processing of digital time series data (or want to learn!). This includes power spectral densities, spectrograms, digital filtering, whitening, audio manipulation. This is a vast and complex set of topics, but we will cover many of the basics in this tutorial.

If you are a beginner, here are some resources from the web:

- http://101science.com/dsp.htm
- https://georgemdallas.wordpress.com/2014/05/14/wavelets-4-dummies-signal-processing-fourier-transforms-and-heisenberg/
- https://en.wikipedia.org/wiki/Signal_processing
- https://en.wikipedia.org/wiki/Spectral_density
- https://en.wikipedia.org/wiki/Spectrogram
- http://greenteapress.com/thinkdsp/
- https://en.wikipedia.org/wiki/Digital_filter

And, well, lots more - google it!

Download the data

- Download the data files from LOSC:
- We will use the hdf5 files, both H1 and L1, with durations of 32 and 4096 seconds around GW150914, sampled at 16384 and 4096 Hz :
  - https://losc.ligo.org/s/events/GW150914/H-H1_LOSC_4_V1-1126259446-32.hdf5
  - https://losc.ligo.org/s/events/GW150914/L-L1_LOSC_4_V1-1126259446-32.hdf5
  - https://losc.ligo.org/s/events/GW150914/H-H1_LOSC_16_V1-1126259446-32.hdf5
  - https://losc.ligo.org/s/events/GW150914/L-L1_LOSC_16_V1-1126259446-32.hdf5
  - https://losc.ligo.org/s/events/GW150914/GW150914_4_NR_waveform.txt
- Download the python functions to read the data: https://losc.ligo.org/s/sample_code/readligo.py
- From a unix/mac-osx command line, you can use wget; for example,
  - wget https://losc.ligo.org/s/events/GW150914/H-H1_LOSC_4_V1-1126257414-4096.hdf5
- Put these files in your current directory / folder. Don't mix any other LOSC data files in this directory, or readligo.py may get confused.

Here,

- "H-H1" means that the data come from the LIGO Hanford Observatory site and the LIGO "H1" datector;
- the "4" means the strain time-series data are (down-)sampled from 16384 Hz to 4096 Hz;
- the "V1" means version 1 of this data release;
- "1126257414-4096" means the data starts at GPS time 1126257414 (Mon Sep 14 09:16:37 GMT 2015), duration 4096 seconds;
  - NOTE: GPS time is number of seconds since Jan 6, 1980 GMT. See http://www.oc.nps.edu/oc2902w/gps/timsys.html or https://losc.ligo.org/gps/
- the filetype "hdf5" means the data are in hdf5 format: https://www.hdfgroup.org/HDF5/

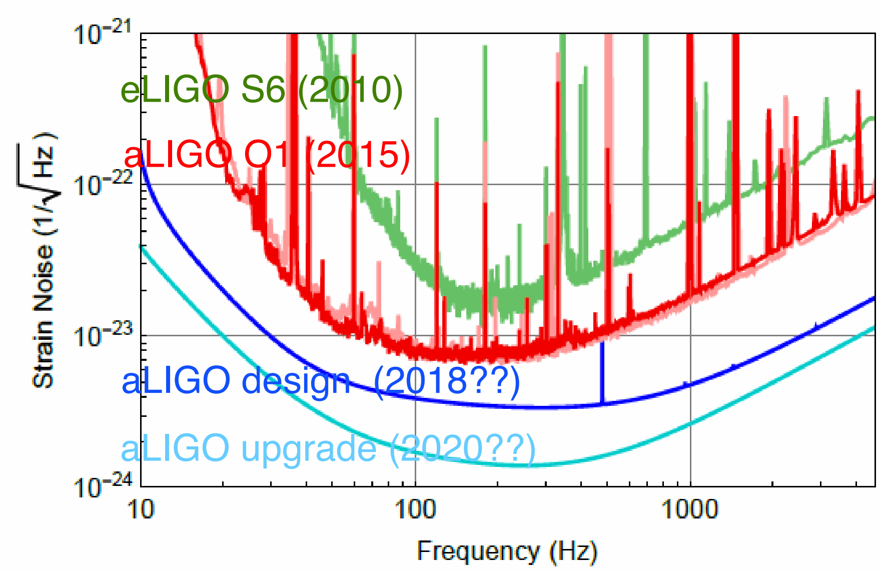
Note that the the 4096 second long files at 16384 Hz sampling rate are fairly big files (125 MB). You won't need them for this tutorial:

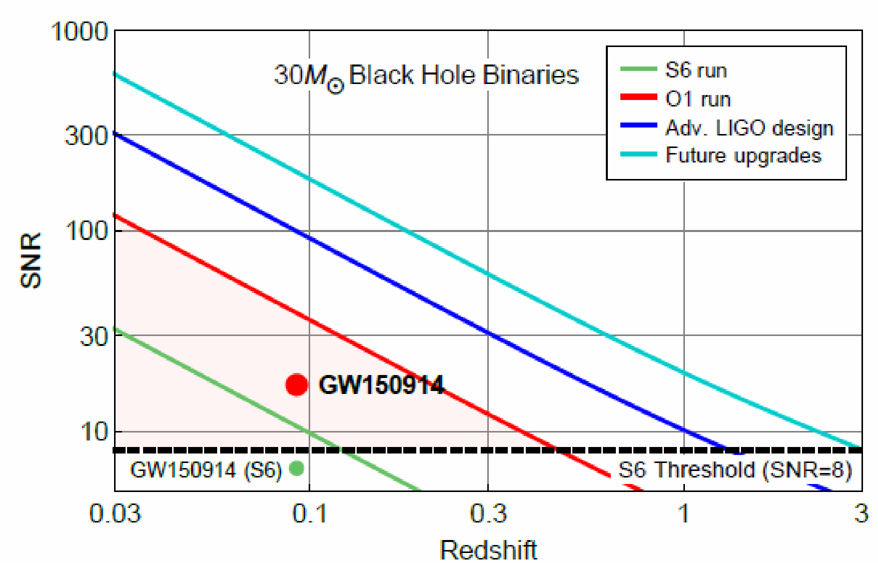
- https://losc.ligo.org/s/events/GW150914/H-H1_LOSC_4_V1-1126257414-4096.hdf5
- https://losc.ligo.org/s/events/GW150914/L-L1_LOSC_4_V1-1126257414-4096.hdf5
- https://losc.ligo.org/s/events/GW150914/H-H1_LOSC_16_V1-1126257414-4096.hdf5

```
# the ipython magic below must be commented out in the .py file, since it doesn't work.
%matplotlib inline
%config InlineBackend.figure_format = 'retina'
import matplotlib.pyplot as plt
```

Caltech LIGO: Jonah Kanner, Roy Williams, AJW

Could it have been found in the initial LIGO detectors?



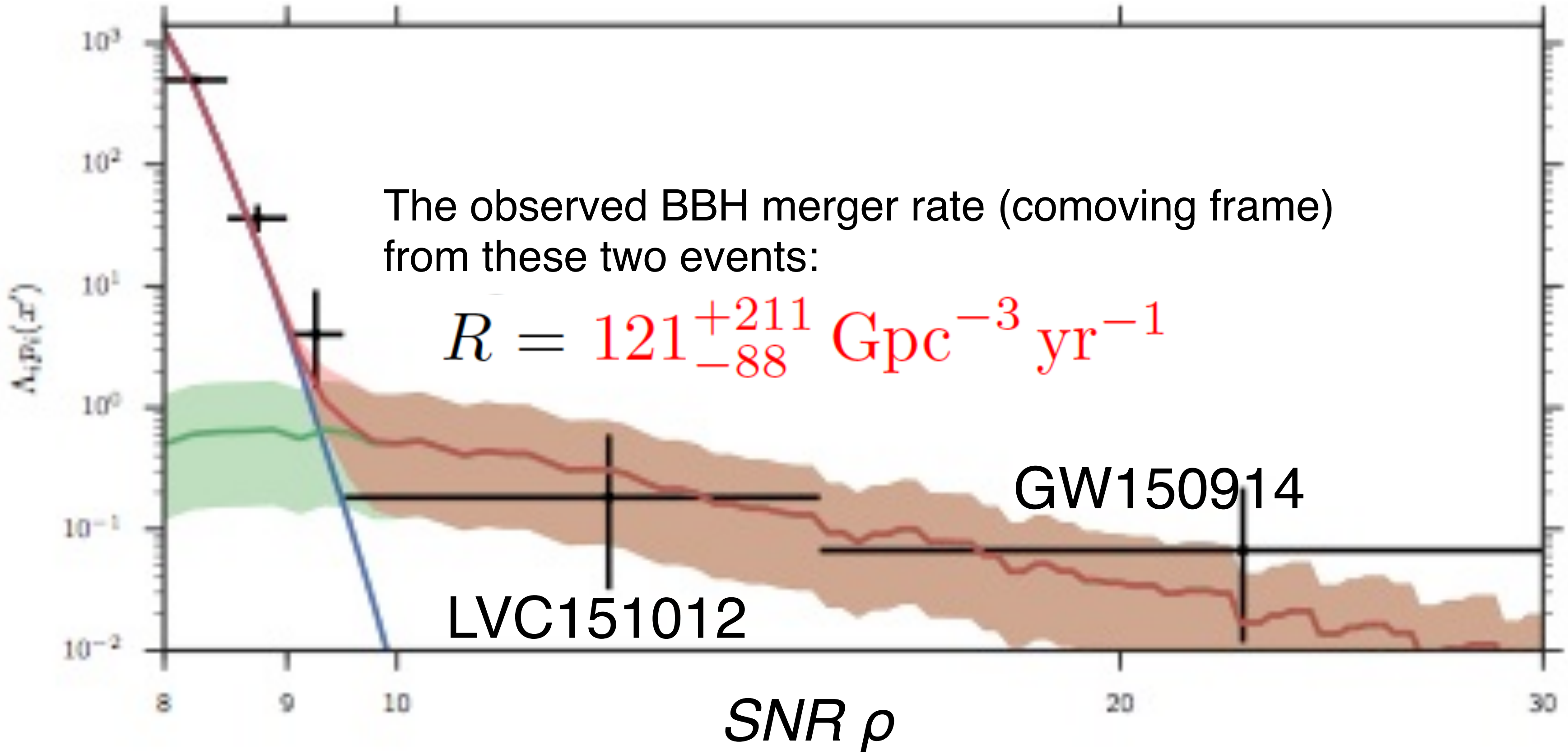


With the S6 detector noise level, this event would have had SNR $\rho \sim 7$,
insufficiently above the noise for detection.
In aLIGO, the signal is very loud!, SNR $\rho \sim 24$.
Aren't sensitive detectors wonderful?

https://dcc.ligo.org/LIGO-P1500237/public/main

Observed BBH merger rate

https://dcc.ligo.org/LIGO-P1500217/public/main



Same ballpark as population synthesis models, CCSN rate, etc

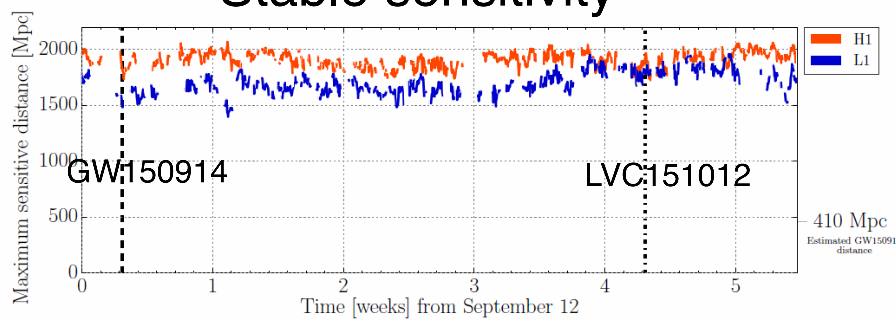
iLIGO+eLIGO BBH rate upper limit: ~< 420 $\mathrm{Gpc}^{-3}\ \mathrm{yr}^{-1}$

Aasi, J. et al. 2013, Phys. Rev. D, 87, 022002, arXiv:1209.6533

Are we SURE it's not a rare instrumental noise fluctuation?

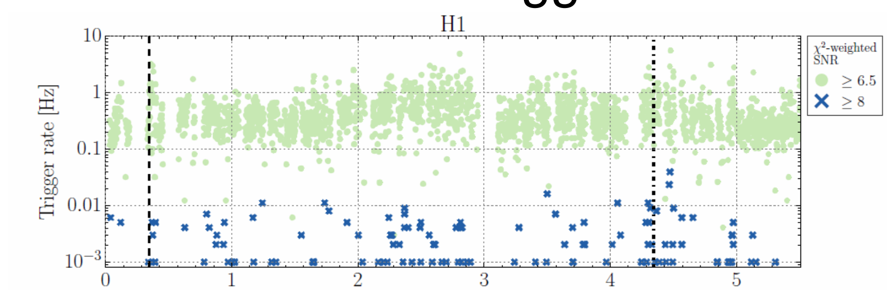
The detectors were behaving quite well, even though it was ER8, 4 days before the official start of O1.

Stable sensitivity



Physical environment monitors (seismometers, magnetometers, microphones, RF monitors, power line monitors, worldwide weather, cosmic ray detectors, etc...) show

Stable noise trigger rate

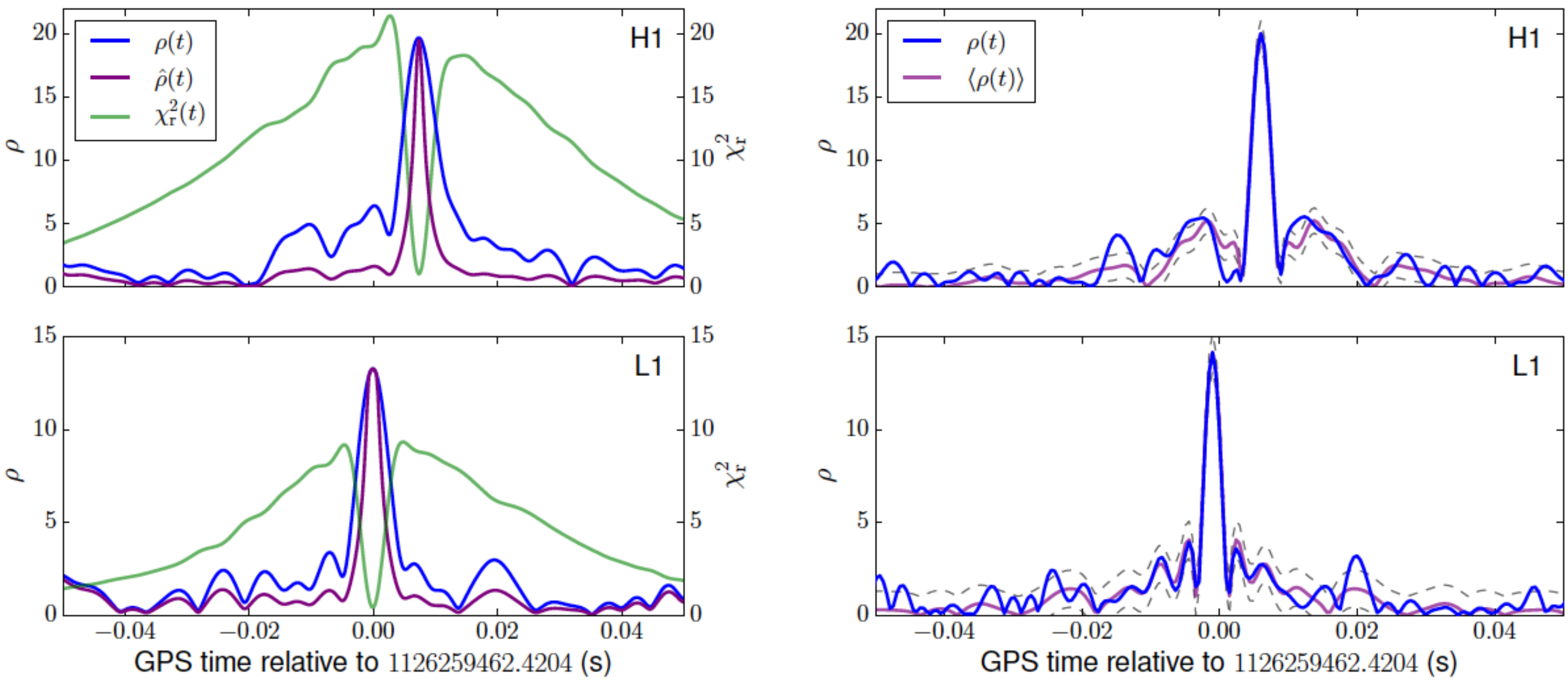


Caltech LIGO: Jessica McIver, Stan Whitcomb, John Zweizig,
Craig Cahillane, Rana Adhikari, AJW

Visitors: Lilli Sun (U Melbourne)

https://dcc.ligo.org/LIGO-P1500238/public/main

SNR time series and χ^2 time series

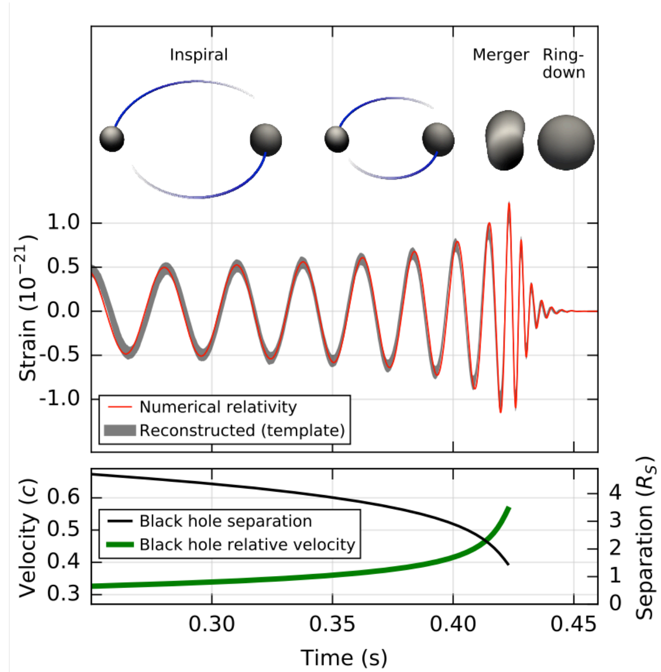


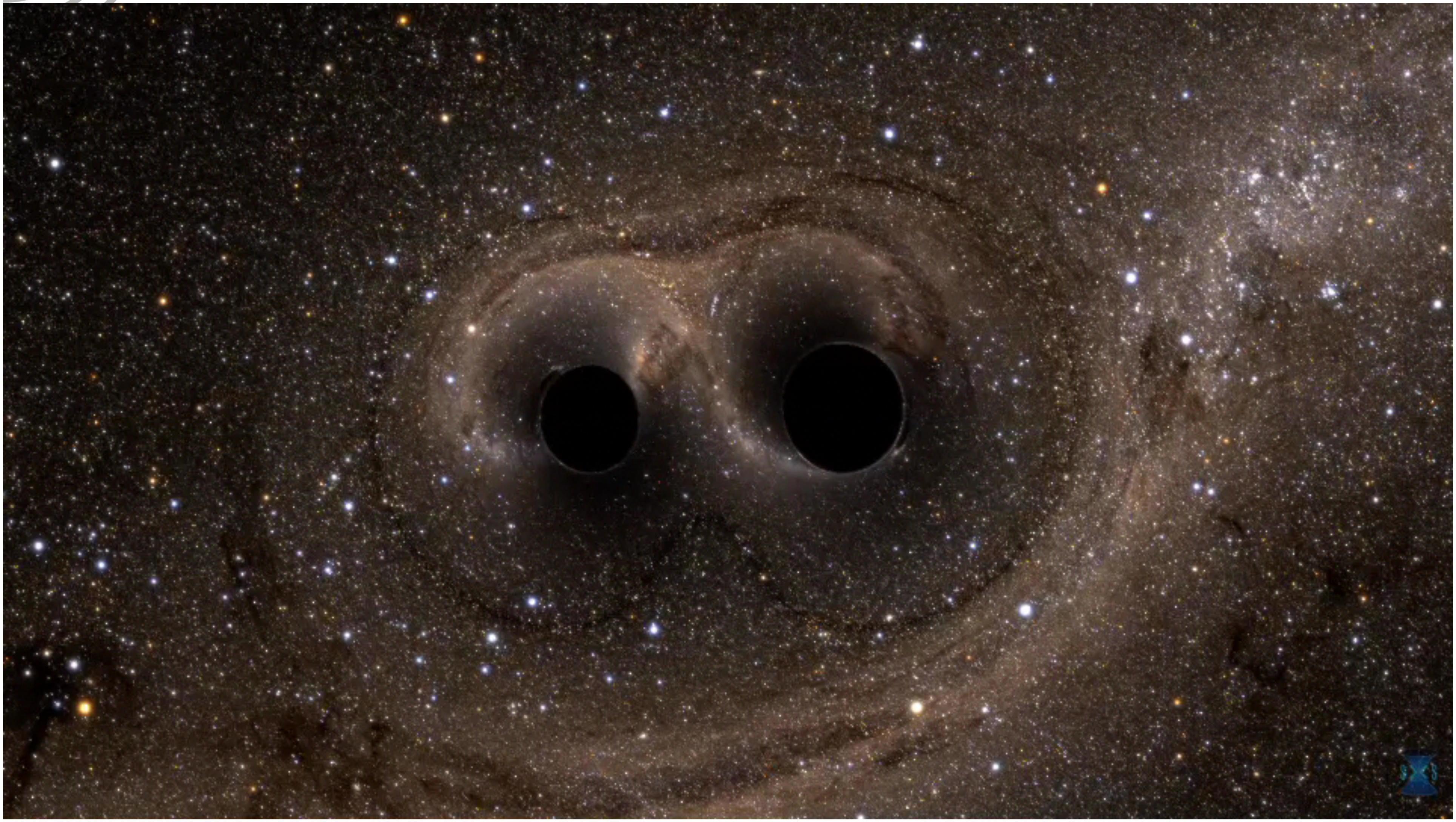
Signal ***ρ(t)*** is consistent with expectation ***<ρ(t)>***

https://dcc.ligo.org/LIGO-P1500238/public/main

What can we learn from one event?

- Excellent **consistency between the observed waveform and the prediction from GR (numerical relativity)** tell us that we are seeing the inspiral of two black holes moving at 0.5c, merging into one BH, which subsequently rings down.
- Such high frequency chirps require extremely **compact orbiting objects of ~ stellar mass**.
- **Black holes (strongly-curved spacetime with event horizons) EXIST**, and emit waves of curved spacetime when perturbed.
  - » Previously, observations of high energy radiation from in-falling matter only told us that compact objects with strong gravity (and perhaps, with event horizons) were present.
- **Binary black holes exist**! Formation scenarios involving common evolution require the binary to survive two core-collapse supernovas.
  Other formation scenarios may be important!
- Two black holes merge into one, which rings down, consistent with **black hole perturbation theory**.
- **GR is tested, for the first time, in the strong (non-linear) and highly dynamical regime**.
- **Masses, spins, sky location, rates, formation mechanisms…**





Numerical relativity (solution to $G_{\mu\nu} = 0$) simulation
(SXS Collaboration, http://www.black-holes.org/)

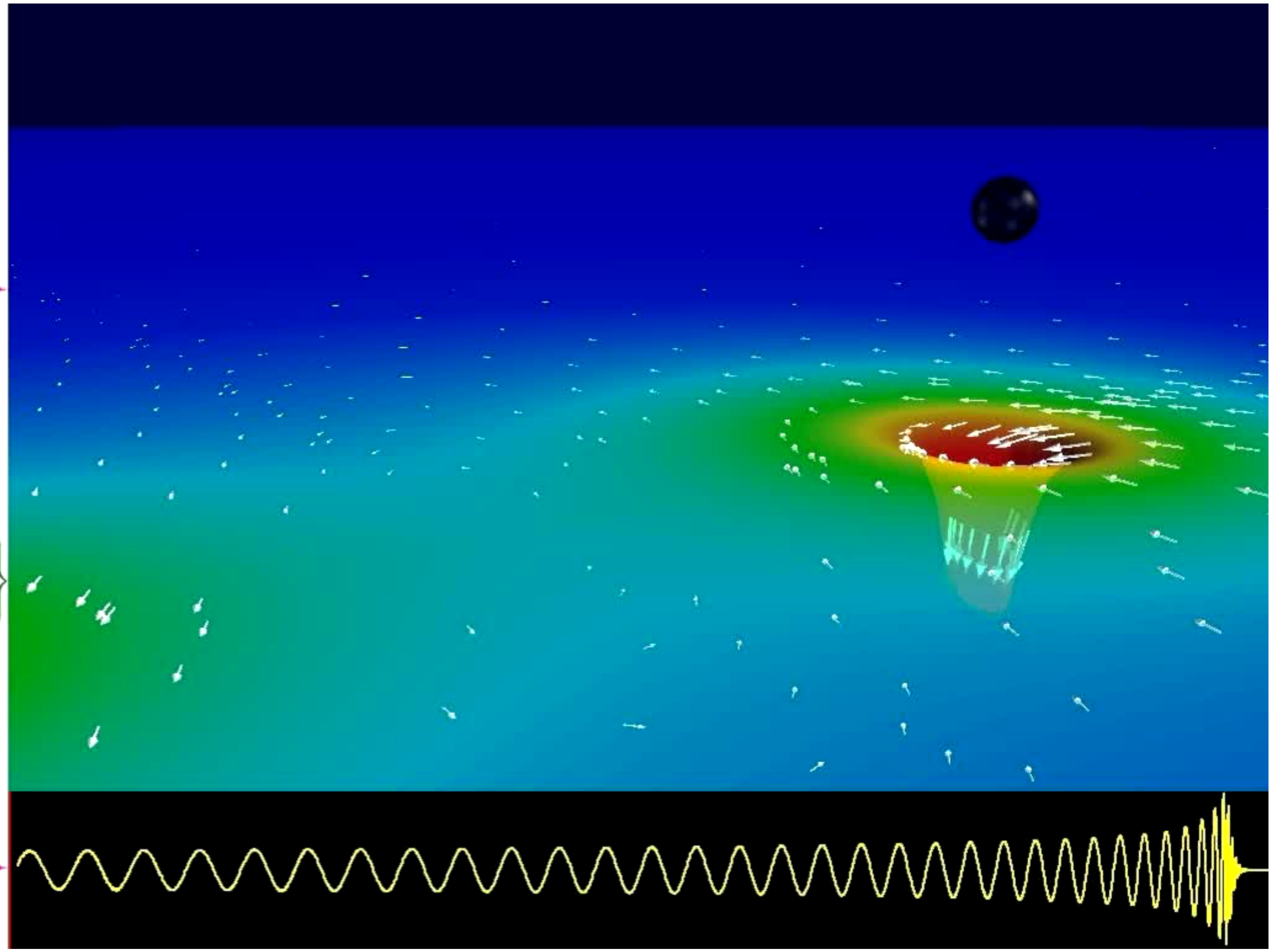
Binary black hole inspiral, merger, ringdown

Binary Black Hole Evolution:
Caltech/Cornell Computer Simulation

Top: 3D view of Black Holes and Orbital Trajectory

Middle: Spacetime curvature:
Depth: Curvature of space
Colors: Rate of flow of time
Arrows: Velocity of flow of space

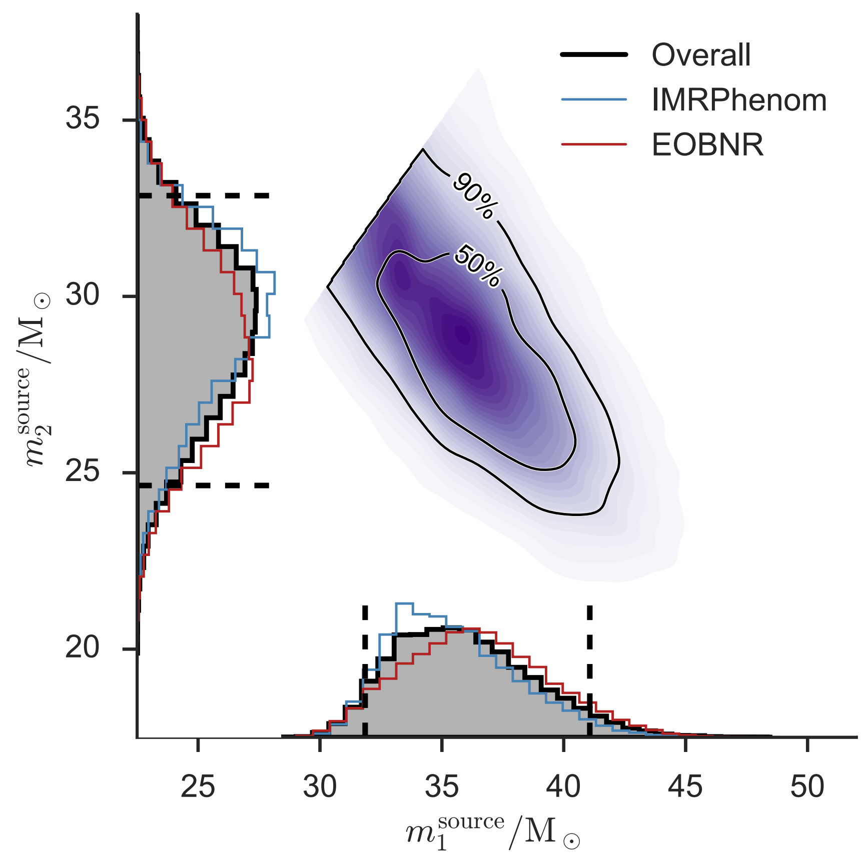
Bottom: Waveform (red line shows current time)



http://www.black-holes.org/explore2.html

Binary Masses

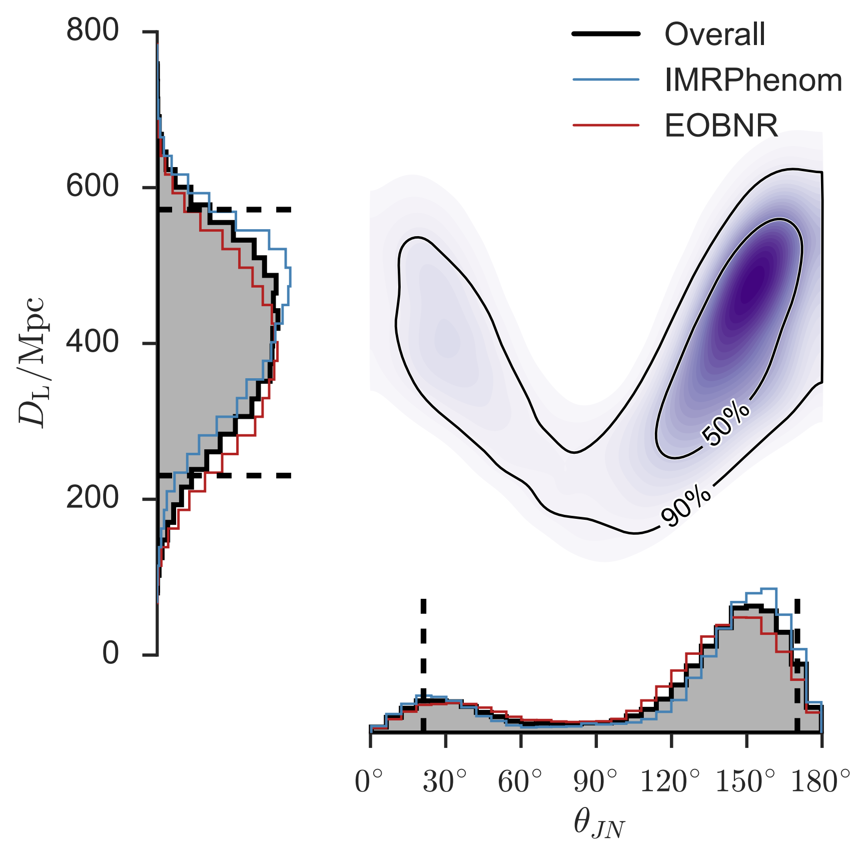
- Measurement of the masses is waveform-model dependent, but for these systems, the waveforms agree well.
- From the inspiral phase evolution $d\phi/dt$, we infer the chirp mass: $\mathcal{M}_c = m_{tot}\, \eta^{3/5}$; $\eta = m_1 m_2 / m^2_{tot}$
- From the merger frequency, we infer the total mass $m_{tot} = m_1 + m_2$.
- In this sweet spot, we measure both well, so we measure m_1 and m_2 reasonably well.



https://dcc.ligo.org/LIGO-P1500218/public/main

Luminosity distance

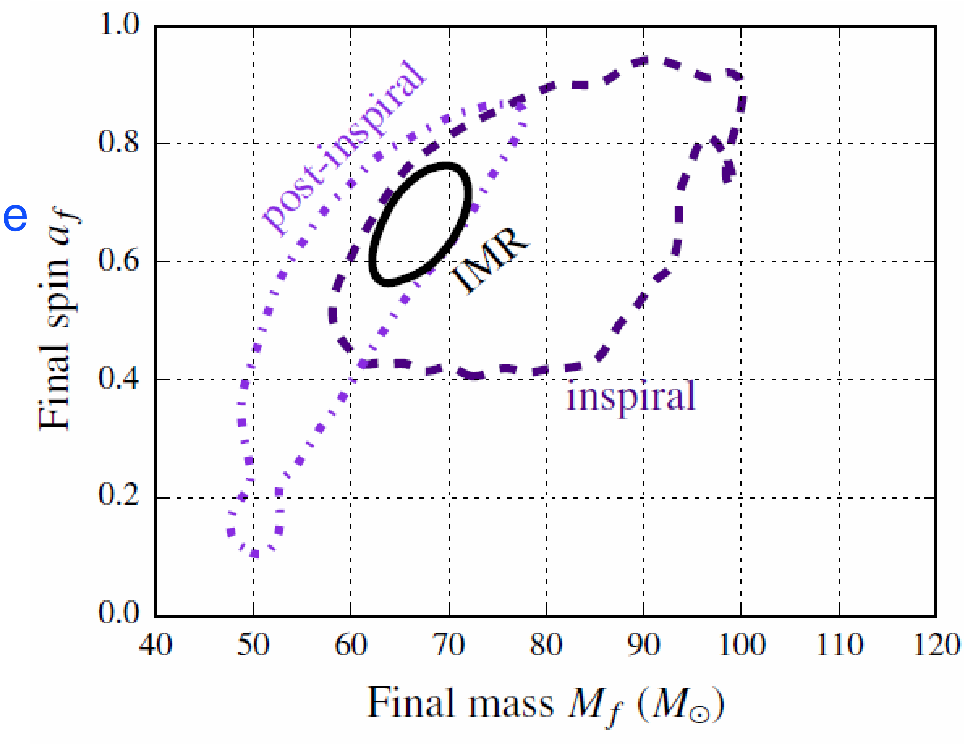
- Compact binary coalescence is a "standard siren", so we can infer the luminosity distance D_L.
- But, it depends on the orientation of the binary orbit wrt line of sight (θ_{JN}): face-on (louder) or edge-on (quieter).
- We can measure this by disentangling the two polarizations (+ and x).
- But this is difficult to do with only two almost-co-aligned detectors.
- Result: strong degeneracy, poor measurement of D_L.
- More detectors (coming!) will help!



https://dcc.ligo.org/LIGO-P1500218/public/main

LIGO Final black hole mass and spin, total emitted energy & luminosity

- SNR of ringdown phase (at f ~ 300 Hz) is not high, so the extraction of the final black hole mass & spin are rather dependent on the model (GR template).
- Nonetheless, we robustly recover $m_{final} \approx m_1 + m_2 - (3\ M_\odot)$ with significant spin.
- $E_{GW} \approx 3\ M_\odot c^2 \approx 5\times10^{54}$ ergs, or ~4.5% of the total mass-energy of the system.
- Roughly 10^{80} gravitons.
- Peak luminosity $L_{GW} \sim 3.6\times10^{54}$ erg/s, briefly outshining the EM energy output of all the stars in the observable universe (by a factor ~ 50).

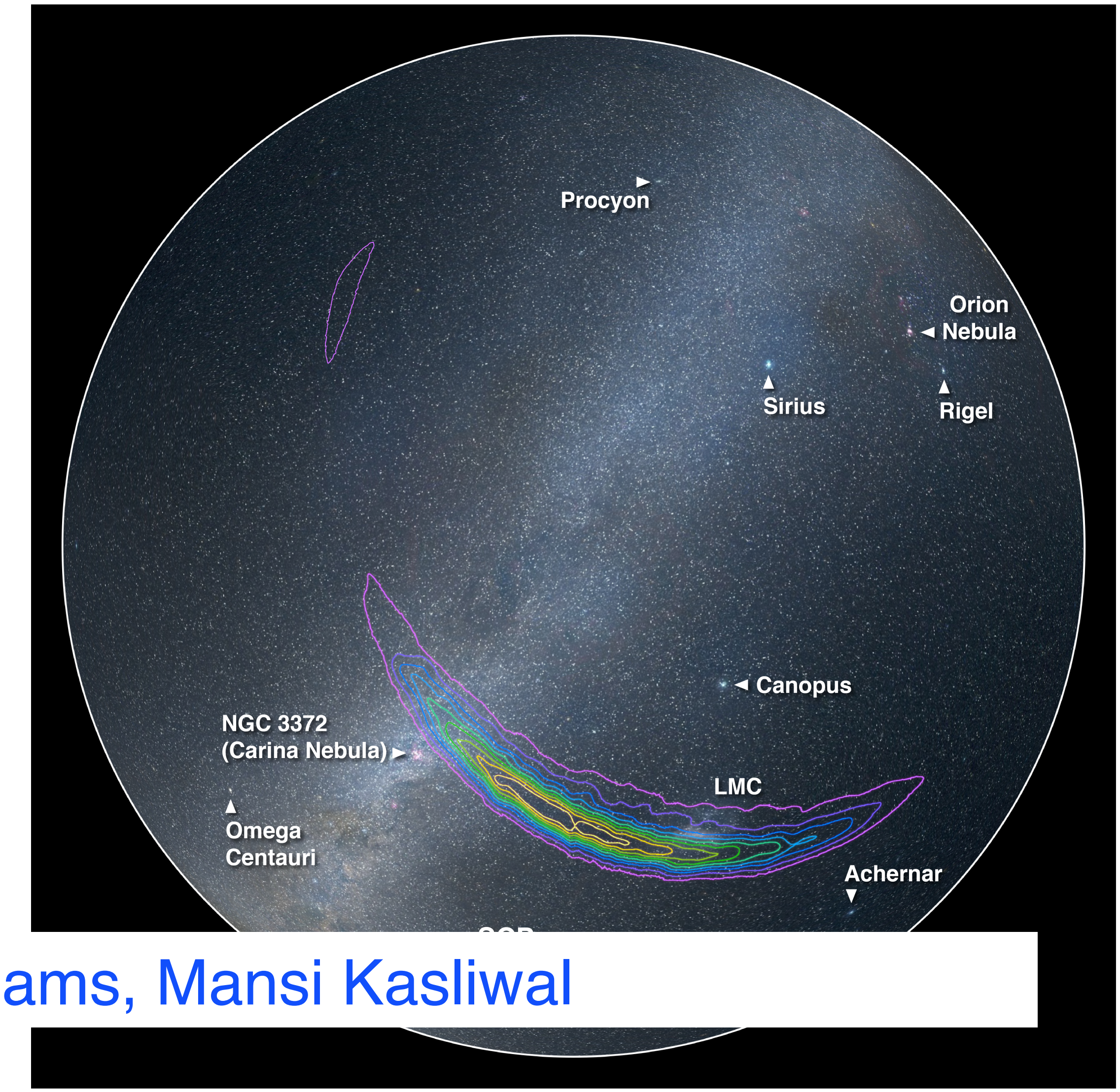


https://dcc.ligo.org/LIGO-P1500218/public/main

Source sky localization

https://dcc.ligo.org/LIGO-P1500227/public/main

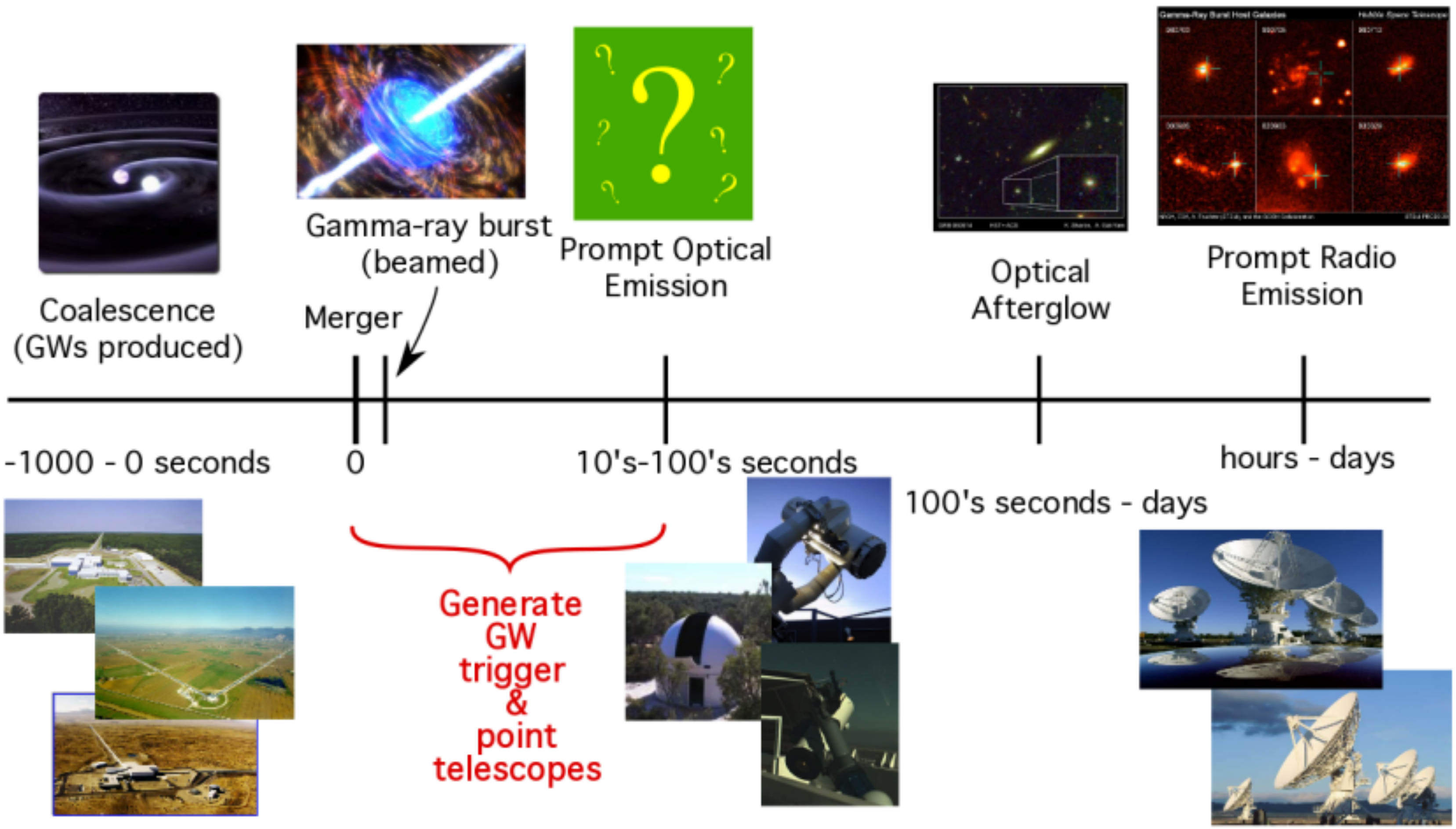
- Localization using timing (7 msec between H1 and L1) information (triangulation), plus amplitude and phase info.
- With only two detectors, localization is poor:
  140 deg^2 at 50% prob,
  590 deg^2 at 90% prob.
- Even though this is a BBH (EM-dim), we alerted partner astronomers, and
  ~20 different instruments imaged this region of the sky!
- Already on arXiv:
  Swift
  Fermi GBM
  DECam



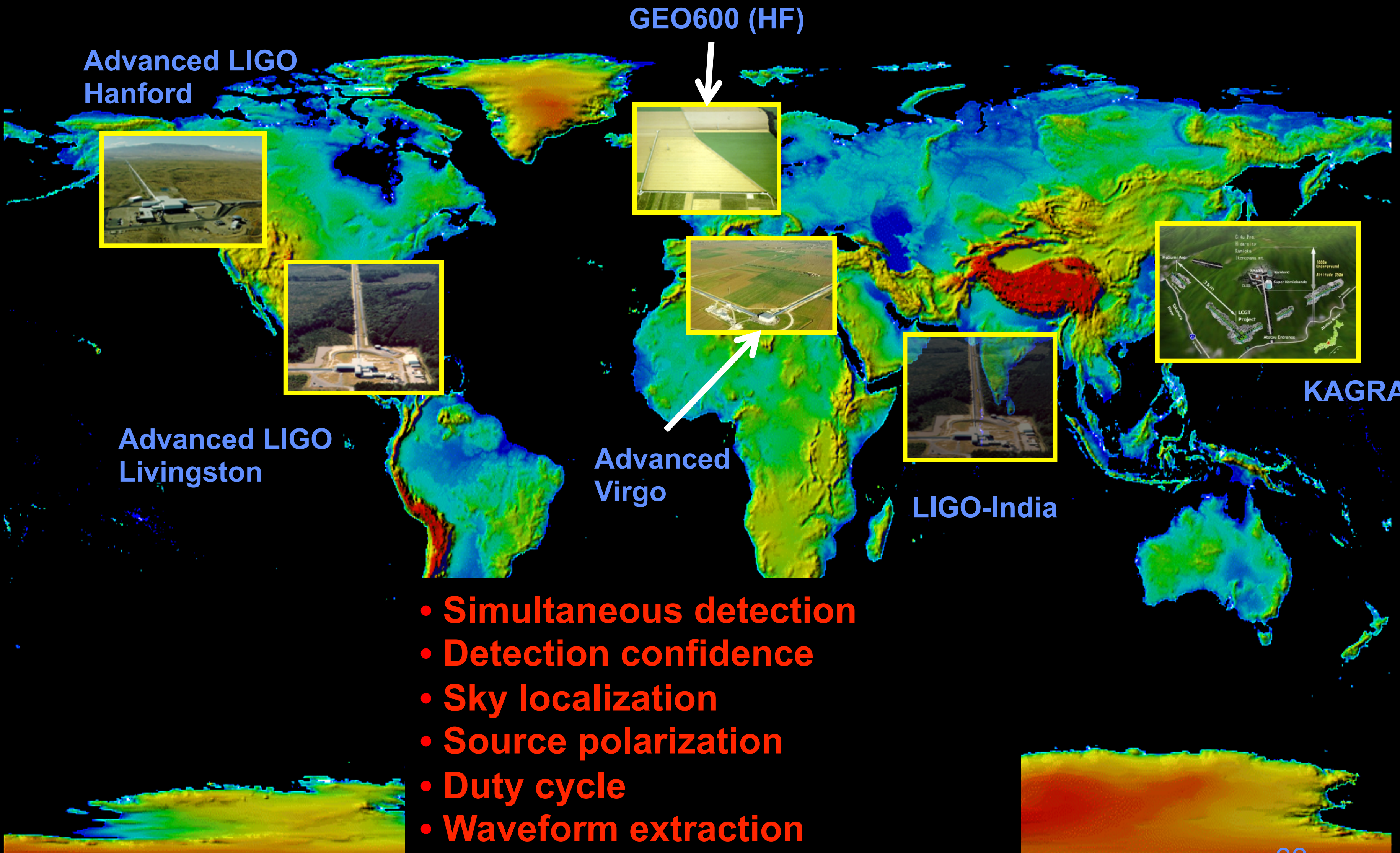
Caltech LIGO: Roy Williams, Mansi Kasliwal

Low-latency identification of transients for rapid (< ~100s) followup

EM counterparts to GW sources (if any) are short-lived and faint



The Advanced GW Detector Network



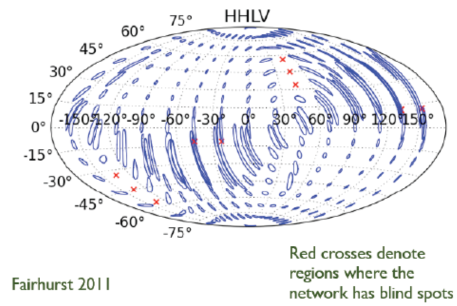
- Simultaneous detection
- Detection confidence
- Sky localization
- Source polarization
- Duty cycle
- Waveform extraction
- Verify light speed propagation

LIGO India is a GO after 5 years of waiting on the Indian government!

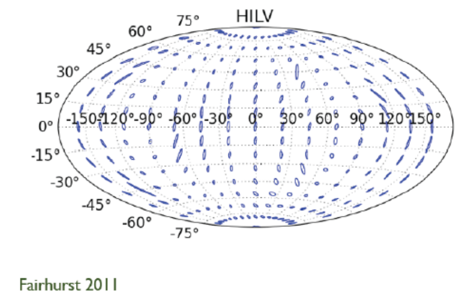
Narendra Modi @narendramodi · Feb 11
Hope to move forward to make even bigger contribution with an advanced gravitational wave detector in the country.
2.1K 4.5K

Narendra Modi @narendramodi · Feb 11
Immensely proud that Indian scientists played an important role in this challenging quest.
1.8K 3.7K

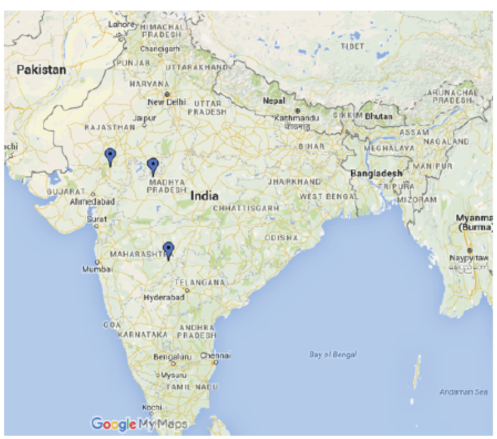
Narendra Modi @narendramodi · Feb 11
Historic detection of gravitational waves opens up new frontier for understanding of universe!
1.9K 3.7K



LIGO+Virgo only



With LIGO-India



Cabinet 17-February, 2016 14:55 IST

Cabinet grants 'in-principle' approval to the LIGO-India mega science proposal

The Union Cabinet chaired by the Prime Minister Shri Narendra Modi has given its 'in principle' approval to the LIGO-India mega science proposal for research on gravitational waves. The proposal, known as LIGO-India project (Laser Interferometer Gravitational-wave Observatory in India) is piloted by Department of Atomic Energy and Department of Science and Technology (DST). The approval coincides with the historic detection of gravitational waves a few days ago that opened up of a new window on the universe to unravel some of its greatest mysteries.

The LIGO-India project will establish a state-of-the-art gravitational wave observatory in India in collaboration with the LIGO Laboratory in the U.S. run by Caltech and MIT.

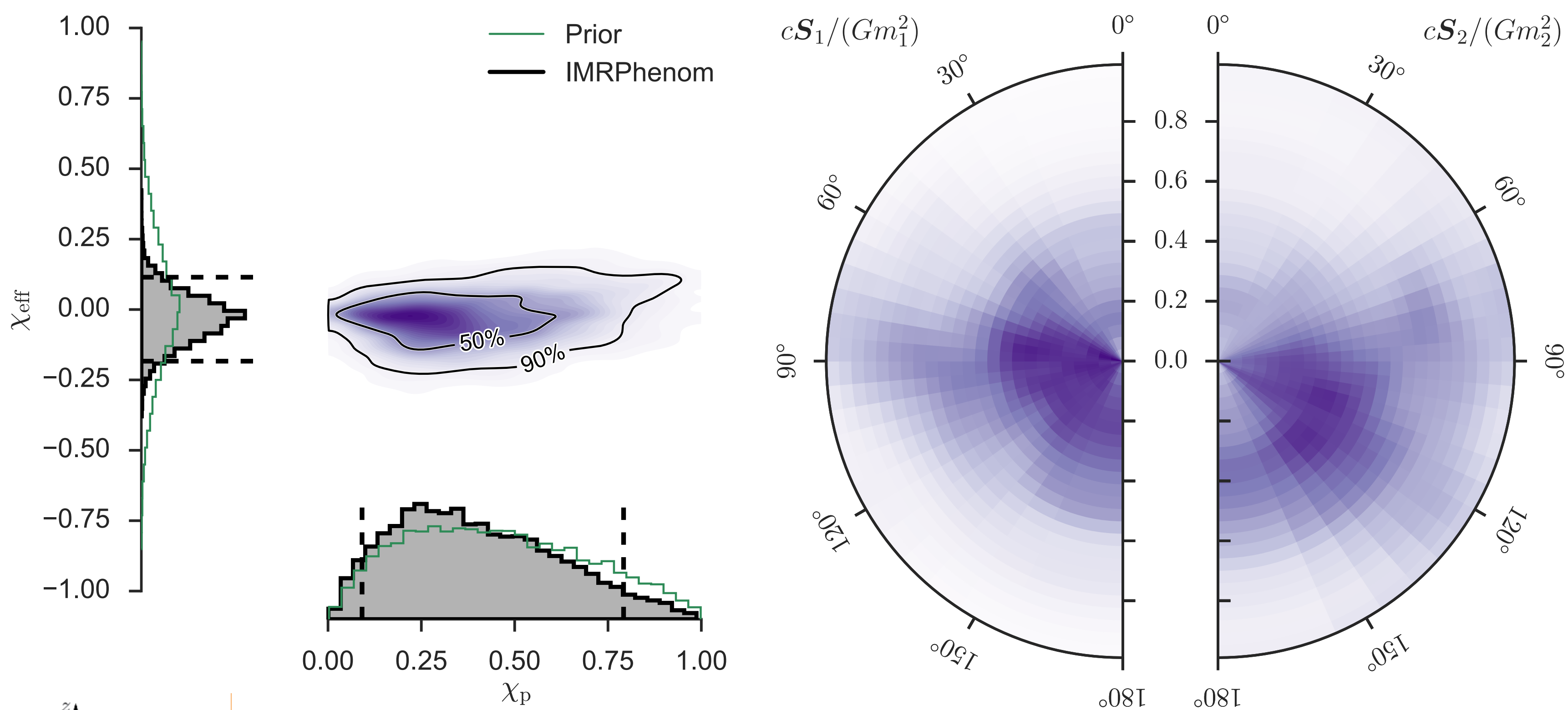
The project will bring unprecedented opportunities for scientists and engineers to dig deeper into the realm of gravitational wave and take global leadership in this new astronomical frontier.

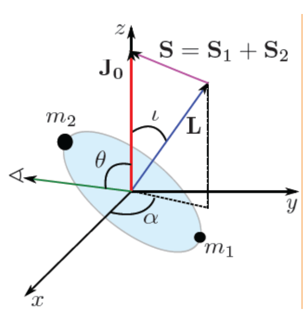
LIGO-India will also bring considerable opportunities in cutting edge technology for the Indian industry which will be engaged in the construction of eight kilometre long beam tube at ultra-high vacuum on a levelled terrain.

The project will motivate Indian students and young scientists to explore newer frontiers of knowledge, and will add further impetus to scientific research in the country.

***

BH spins – aligned with orbital angular momentum, and precessing spin





Caltech LIGO: Patricia Schmidt, Rory Smith
SXS Collaboration

https://dcc.ligo.org/LIGO-P1500218/public/main

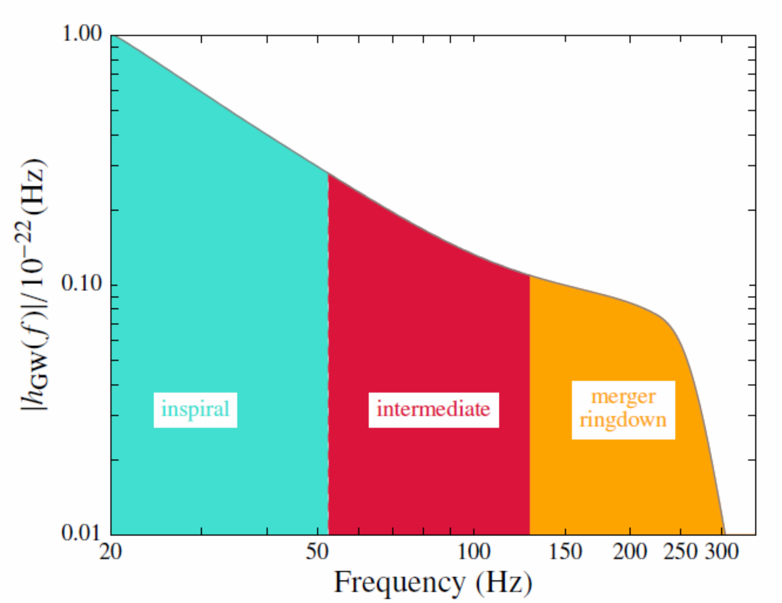
Parameters of two loudest events

https://dcc.ligo.org/LIGO-P1500218/public/main

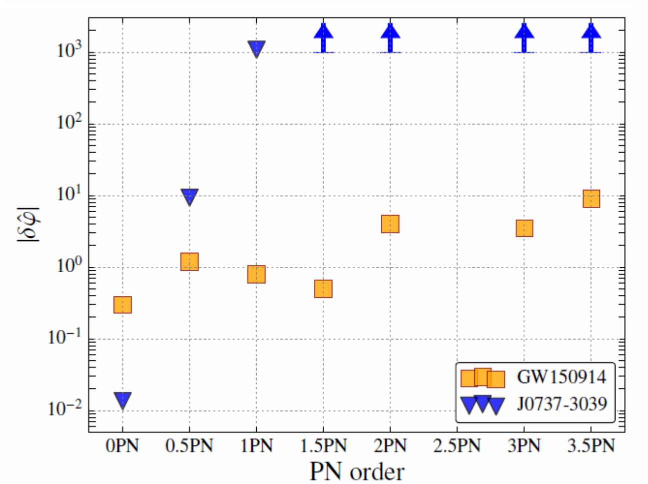
Event	Time (UTC)	FAR (yr^{-1})	$\mathscr{F}$	$\mathscr{M}$ ($M_\odot$)	m_1 ($M_\odot$)	m_2 ($M_\odot$)	χ_{eff}	D_L (Mpc)
GW150914	14 September 2015 09:50:45	$< 5 \times 10^{-6}$	$< 2 \times 10^{-7}$ ($> 5.1\,\sigma$)	28^{+2}_{-2}	36^{+5}_{-4}	29^{+4}_{-4}	$-0.06^{+0.17}_{-0.18}$	410^{+160}_{-180}
LVT151012	12 October 2015 09:54:43	0.44	0.02 ($2.1\,\sigma$)	15^{+1}_{-1}	23^{+18}_{-5}	13^{+4}_{-5}	$0.0^{+0.3}_{-0.2}$	1100^{+500}_{-500}

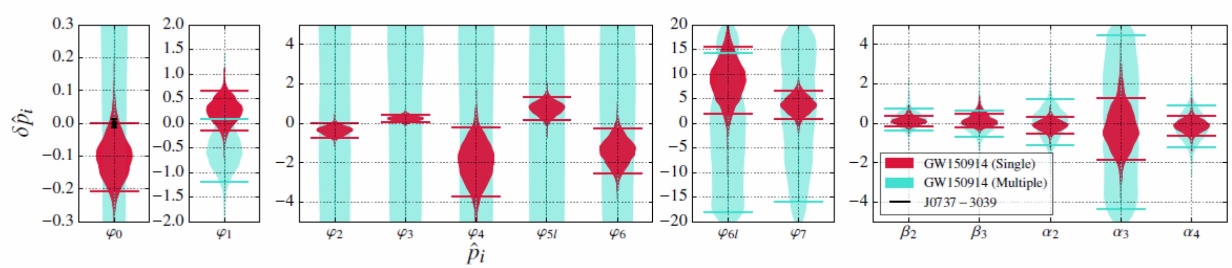
	EOBNR	IMRPhenom	Overall
Detector-frame total mass $M/M_\odot$	$70.3^{+5.3}_{-4.8}$	$70.7^{+3.8}_{-4.0}$	$70.5^{+4.6\pm0.9}_{-4.5\pm1.0}$
Detector-frame chirp mass $\mathcal{M}/M_\odot$	$30.2^{+2.5}_{-1.9}$	$30.5^{+1.7}_{-1.8}$	$30.3^{+2.1\pm0.4}_{-1.9\pm0.4}$
Detector-frame primary mass $m_1/M_\odot$	$39.4^{+5.5}_{-4.9}$	$38.3^{+5.5}_{-3.5}$	$38.8^{+5.6\pm0.9}_{-4.1\pm0.3}$
Detector-frame secondary mass $m_2/M_\odot$	$30.9^{+4.8}_{-4.4}$	$32.2^{+3.6}_{-5.0}$	$31.6^{+4.2\pm0.1}_{-4.9\pm0.6}$
Detector-frame final mass $M_f/M_\odot$	$67.1^{+4.6}_{-4.4}$	$67.4^{+3.4}_{-3.6}$	$67.3^{+4.1\pm0.8}_{-4.0\pm0.9}$
Source-frame total mass $M^{source}/M_\odot$	$65.0^{+5.0}_{-4.4}$	$64.6^{+4.1}_{-3.5}$	$64.8^{+4.6\pm1.0}_{-3.9\pm0.5}$
Source-frame chirp mass $\mathcal{M}^{source}/M_\odot$	$27.9^{+2.3}_{-1.8}$	$27.9^{+1.8}_{-1.6}$	$27.9^{+2.1\pm0.4}_{-1.7\pm0.2}$
Source-frame primary mass $m_1^{source}/M_\odot$	$36.3^{+5.3}_{-4.5}$	$35.1^{+5.2}_{-3.3}$	$35.7^{+5.4\pm1.1}_{-3.8\pm0.0}$
Source-frame secondary mass $m_2^{source}/M_\odot$	$28.6^{+4.4}_{-4.2}$	$29.5^{+3.3}_{-4.5}$	$29.1^{+3.8\pm0.2}_{-4.4\pm0.5}$
Source-fame final mass $M_f^{source}/M_\odot$	$62.0^{+4.4}_{-4.0}$	$61.6^{+3.7}_{-3.1}$	$61.8^{+4.2\pm0.9}_{-3.5\pm0.4}$
Mass ratio q	$0.79^{+0.18}_{-0.19}$	$0.84^{+0.14}_{-0.21}$	$0.82^{+0.16\pm0.01}_{-0.21\pm0.03}$
Effective inspiral spin parameter χ_{eff}	$-0.09^{+0.19}_{-0.17}$	$-0.03^{+0.14}_{-0.15}$	$-0.06^{+0.17\pm0.01}_{-0.18\pm0.07}$
Dimensionless primary spin magnitude a_1	$0.32^{+0.45}_{-0.28}$	$0.31^{+0.51}_{-0.27}$	$0.31^{+0.48\pm0.04}_{-0.28\pm0.01}$
Dimensionless secondary spin magnitude a_2	$0.57^{+0.40}_{-0.51}$	$0.39^{+0.50}_{-0.34}$	$0.46^{+0.48\pm0.07}_{-0.42\pm0.01}$
Final spin a_f	$0.67^{+0.06}_{-0.08}$	$0.67^{+0.05}_{-0.05}$	$0.67^{+0.05\pm0.00}_{-0.07\pm0.03}$
Luminosity distance D_L/Mpc	390^{+170}_{-180}	440^{+140}_{-180}	$410^{+160\pm20}_{-180\pm40}$
Source redshift z	$0.083^{+0.033}_{-0.036}$	$0.093^{+0.028}_{-0.036}$	$0.088^{+0.031\pm0.004}_{-0.038\pm0.009}$
Upper bound on primary spin magnitude a_1	0.65	0.71	0.69 ± 0.05
Upper bound on secondary spin magnitude a_2	0.93	0.81	0.88 ± 0.10
Lower bound on mass ratio q	0.64	0.67	0.65 ± 0.03
Log Bayes factor $\ln \mathcal{B}_{s/n}$	288.7 ± 0.2	290.1 ± 0.2	—

Tests of consistency with predictions from General Relativity



waveform regime	parameter	f−dependence	median single	median multiple	GR quantile single	GR quantile multiple	$\log_{10} B^{GR}_{model}$ single	$\log_{10} B^{GR}_{model}$ multiple
early-inspiral regime	$\delta\hat{\varphi}_0$	$f^{-5/3}$	$-0.1^{+0.1}_{-0.1}$	$1.3^{+3.0}_{-3.2}$	0.94	0.30	1.9 ± 0.2	3.7 ± 0.6
	$\delta\hat{\varphi}_1$	$f^{-4/3}$	$0.3^{+0.4}_{-0.4}$	$-0.5^{+0.6}_{-0.6}$	0.16	0.93	1.6 ± 0.2	
	$\delta\hat{\varphi}_2$	f^{-1}	$-0.4^{+0.3}_{-0.4}$	$-1.6^{+18.8}_{-16.6}$	0.96	0.56	1.2 ± 0.2	
	$\delta\hat{\varphi}_3$	$f^{-2/3}$	$0.2^{+0.2}_{-0.2}$	$2.0^{+13.4}_{-13.9}$	0.02	0.42	1.2 ± 0.2	
	$\delta\hat{\varphi}_4$	$f^{-1/3}$	$-1.9^{+1.6}_{-1.7}$	$-1.9^{+19.3}_{-16.4}$	0.98	0.56	0.3 ± 0.2	
	$\delta\hat{\varphi}_{5l}$	$\log(f)$	$0.8^{+0.5}_{-0.6}$	$-1.4^{+18.6}_{-16.9}$	0.01	0.55	0.7 ± 0.4	
	$\delta\hat{\varphi}_6$	$f^{1/3}$	$-1.4^{+1.1}_{-1.1}$	$1.2^{+16.8}_{-18.9}$	0.99	0.47	0.4 ± 0.2	
	$\delta\hat{\varphi}_{6l}$	$f^{1/3}\log(f)$	$8.9^{+6.8}_{-6.8}$	$-1.9^{+19.1}_{-16.1}$	0.02	0.57	-0.3 ± 0.2	
	$\delta\hat{\varphi}_7$	$f^{2/3}$	$3.8^{+2.9}_{-2.9}$	$3.2^{+15.1}_{-19.2}$	0.02	0.41	-0.0 ± 0.2	
intermediate regime	$\delta\hat{\beta}_2$	$\log f$	$0.1^{+0.4}_{-0.3}$	$0.2^{+0.6}_{-0.5}$	0.24	0.28	1.4 ± 0.2	2.3 ± 0.2
	$\delta\hat{\beta}_3$	f^{-3}	$0.1^{+0.6}_{-0.3}$	$-0.0^{+0.8}_{-0.7}$	0.31	0.56	1.2 ± 0.4	
merger-ringdown regime	$\delta\hat{\alpha}_2$	f^{-1}	$-0.1^{+0.4}_{-0.4}$	$0.0^{+1.0}_{-1.2}$	0.68	0.50	1.2 ± 0.2	2.1 ± 0.4
	$\delta\hat{\alpha}_3$	$f^{3/4}$	$-0.3^{+1.9}_{-1.5}$	$0.0^{+4.4}_{-4.4}$	0.60	0.51	0.7 ± 0.2	
	$\delta\hat{\alpha}_4$	$\tan^{-1}(af+b)$	$-0.1^{+0.5}_{-0.5}$	$-0.1^{+1.1}_{-1.0}$	0.68	0.62	1.1 ± 0.2	





https://dcc.ligo.org/LIGO-P1500213/public/main

Mass of the graviton

A graviton mass

$$E^2 = p^2c^2 + m_g^2c^4$$

and associated Compton wavelength

$$\lambda_g = h/(m_g c)$$

results in frequency-dependent velocity

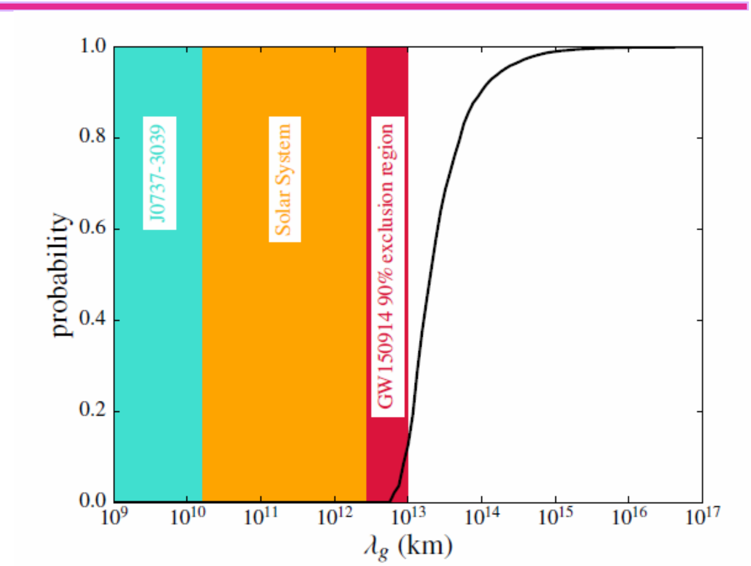
$$v_g^2/c^2 \equiv c^2p^2/E^2 = 1 - h^2c^2/(\lambda_g^2 E^2)$$

and dispersion causes distortion of the phase evolution of the waveform (wrt massless theory)

$$\Phi_{\mathrm{MG}}(f) = -(\pi D c)/[\lambda_g^2(1+z)f]$$

Agreement of observed waveform with theory allows us to set the bound:

$$m_g \leq 1.2 \times 10^{-22}\ \mathrm{eV/c^2}\ \text{at 90\% confidence}$$

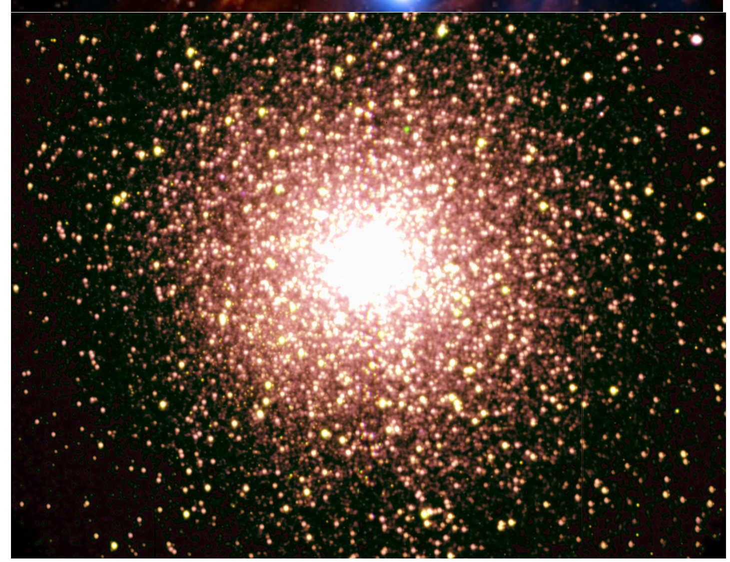


$\lambda_g > 10^{13}$ km

https://dcc.ligo.org/LIGO-P1500213/public/main

Formation mechanisms

- How do massive binary black hole systems form?
- Common envelope evolution of isolated binaries: two massive stars survive successive CCSNe
- Dynamical capture of isolated black holes in N-body exchange interactions.
- Even the most massive stars (60-100 $M_\odot$) can only produce black holes with mass > 20 $M_\odot$ only in low-metalicity environments ($\sim 0.1\ Z_\odot$).
- https://dcc.ligo.org/LIGO-P1500262/public/main



Formation channels

https://dcc.ligo.org/LIGO-P1500262/public/main

Isolated binary

“ordinary”
or
rapid
rotation
or
PopIII

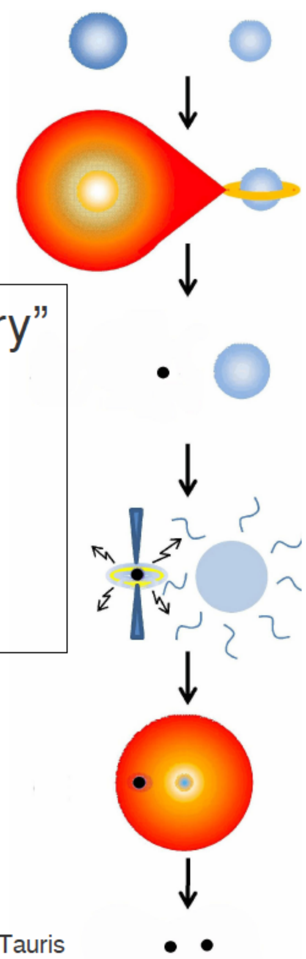


Fig. after Tauris

Dynamical formation

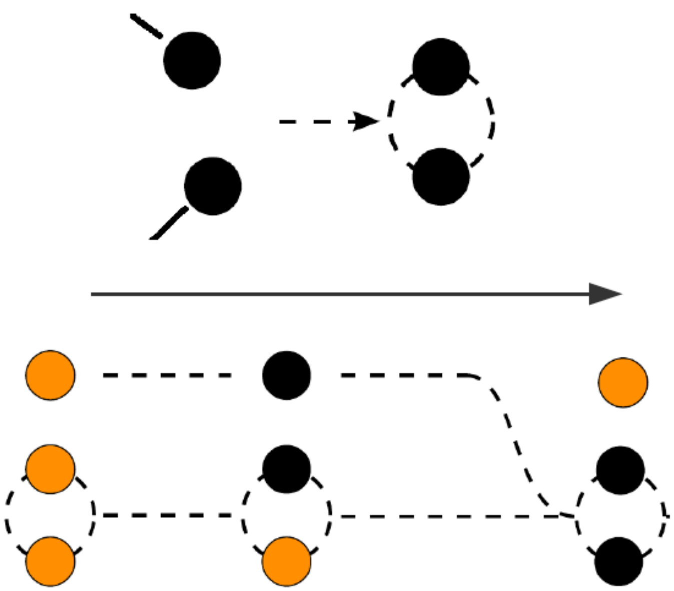
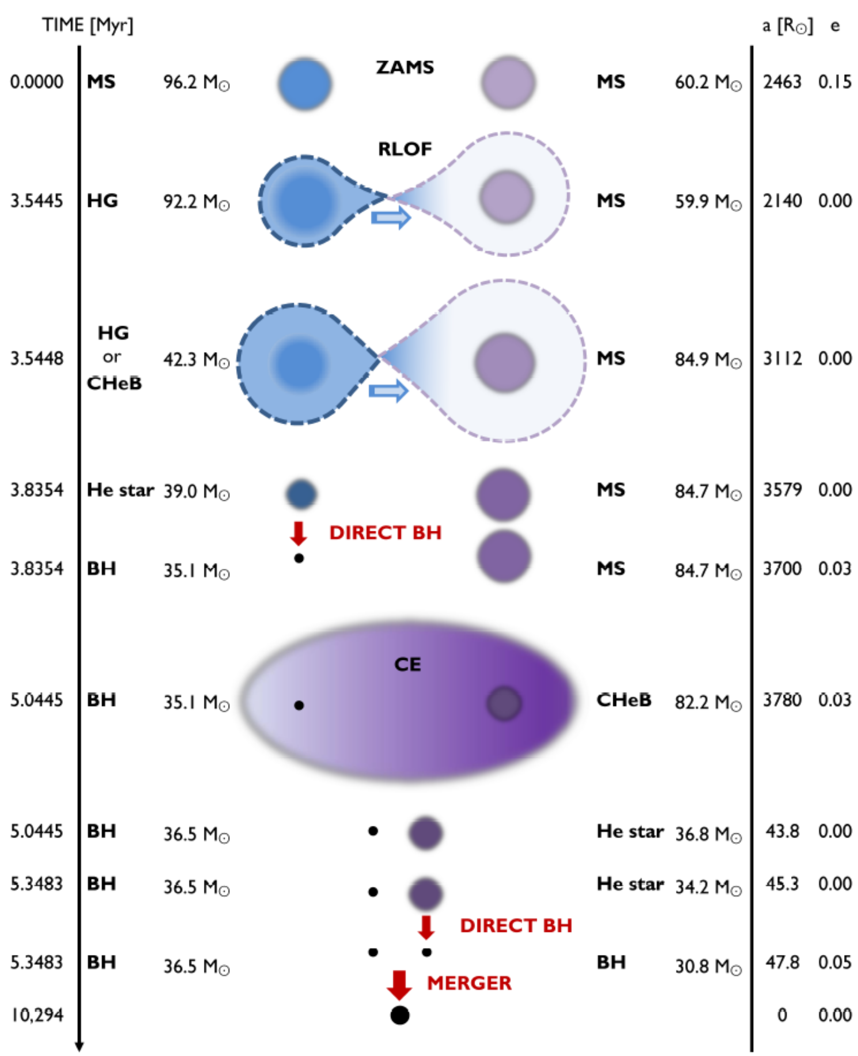


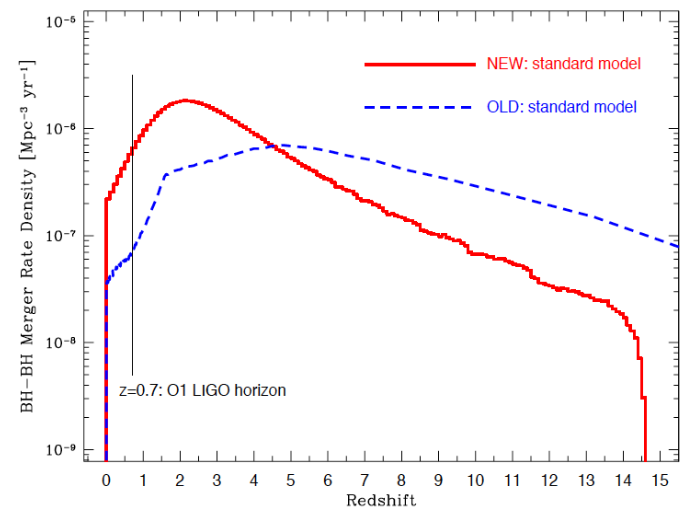
Fig. after Ziosi

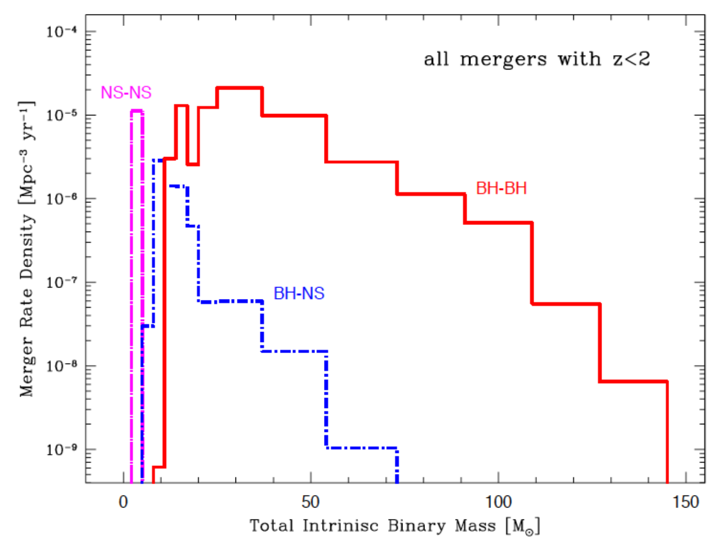
Globular/young clusters/gal. nuclei

“Classic” evolutionary scenario

Belczynski et al, arXiv:1602.04531

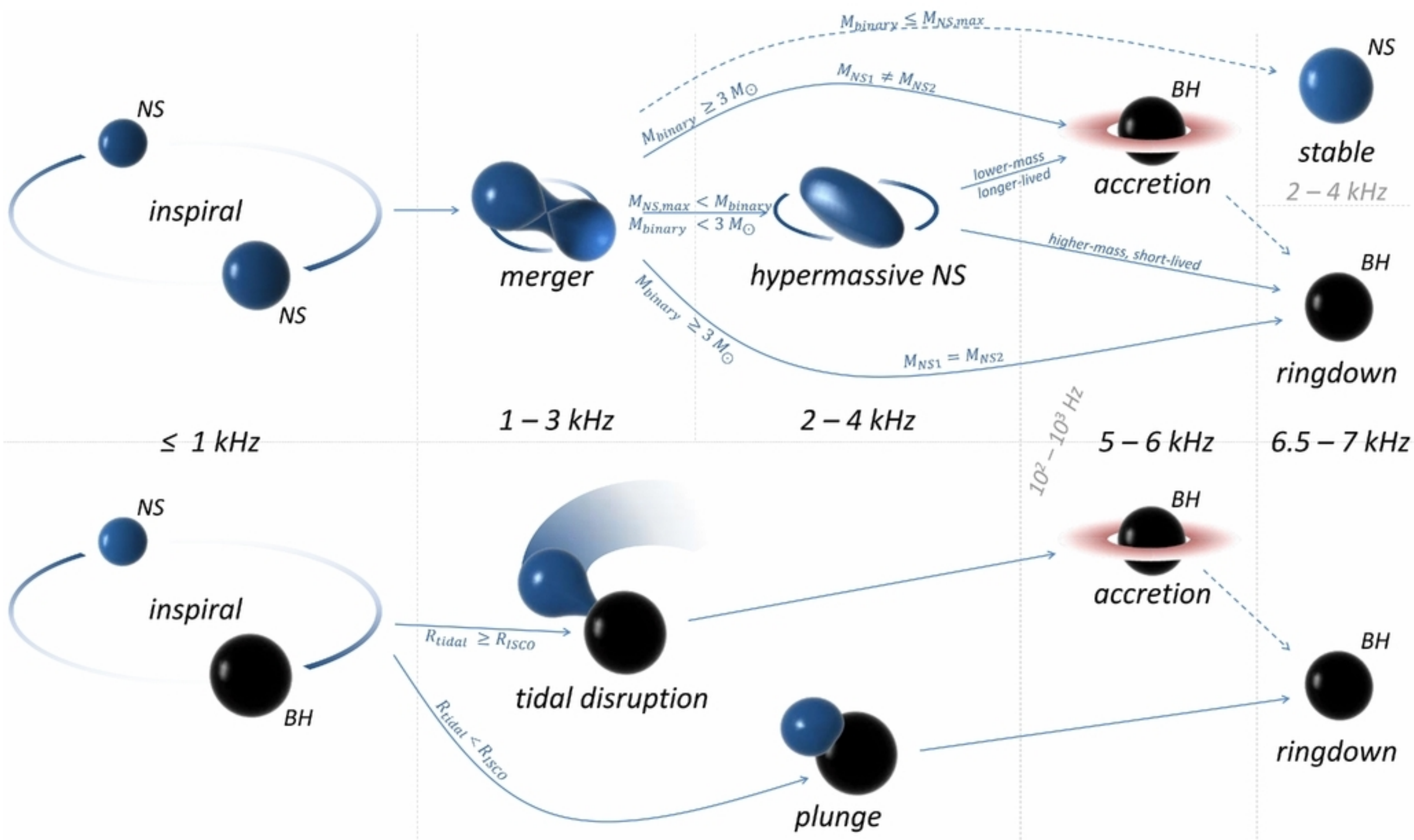






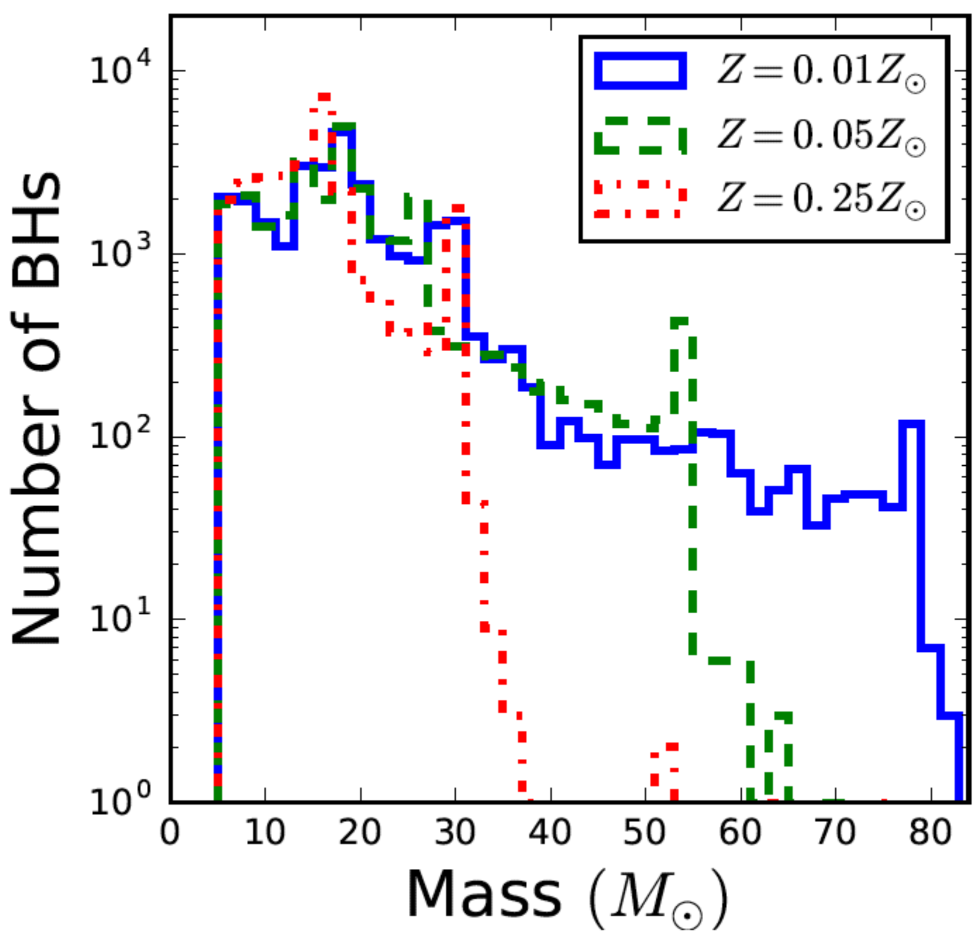
And in the end … binary mergers

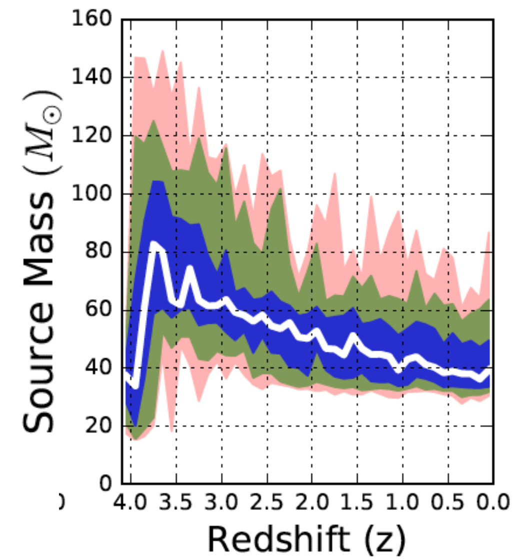
Figure 1 from I Bartos et al 2013 Class. Quantum Grav. 30 123001

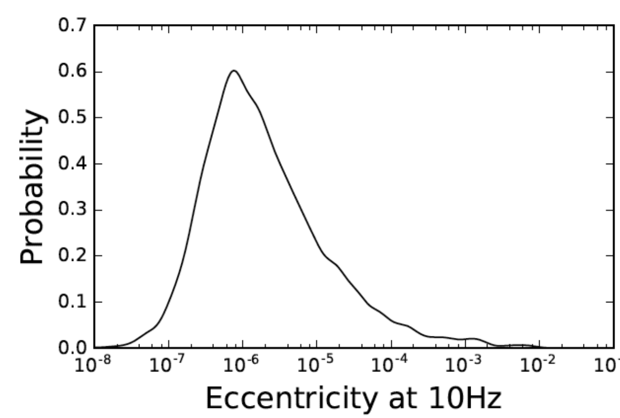


BBH Mergers from Globular Clusters

C. Rodriguez et al, arXiv:1602.02444



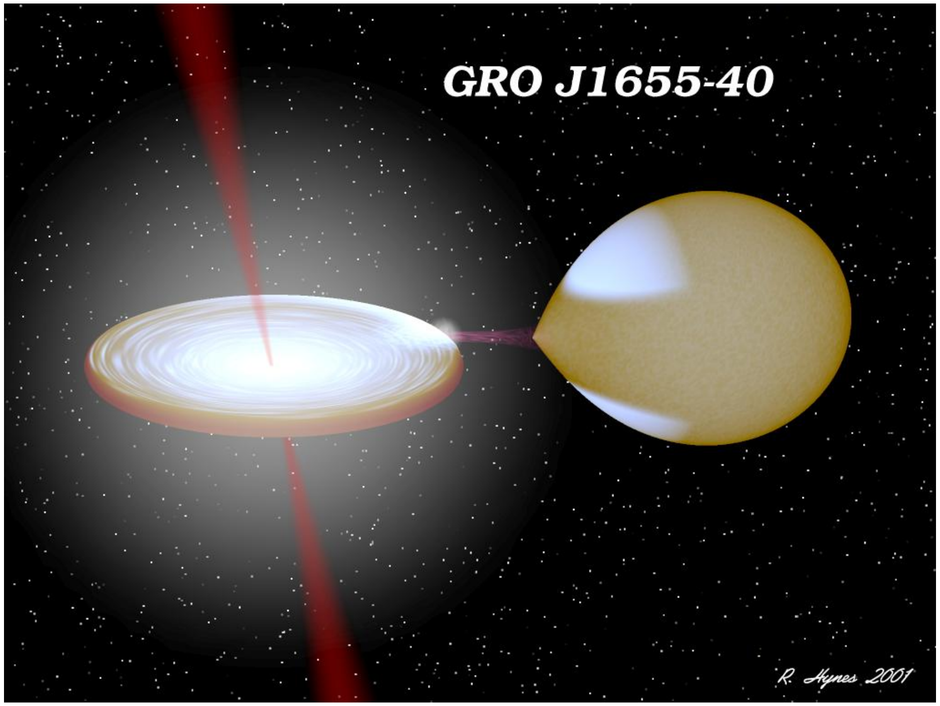


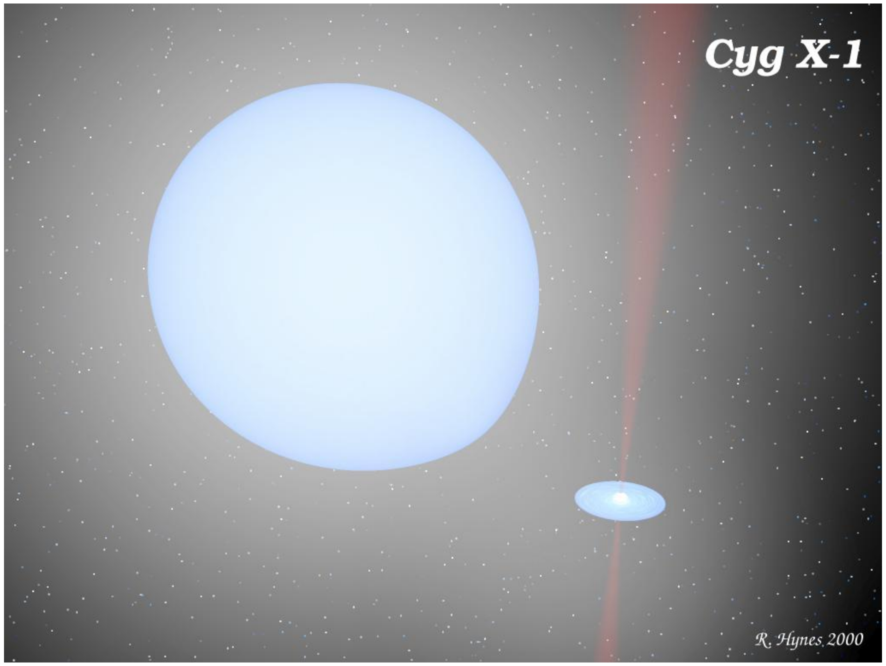


Progenitors of compact binaries

- LMXB

HMXB



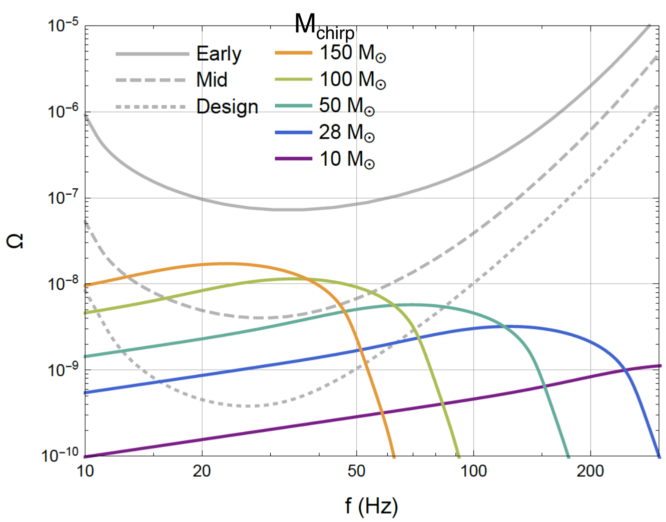


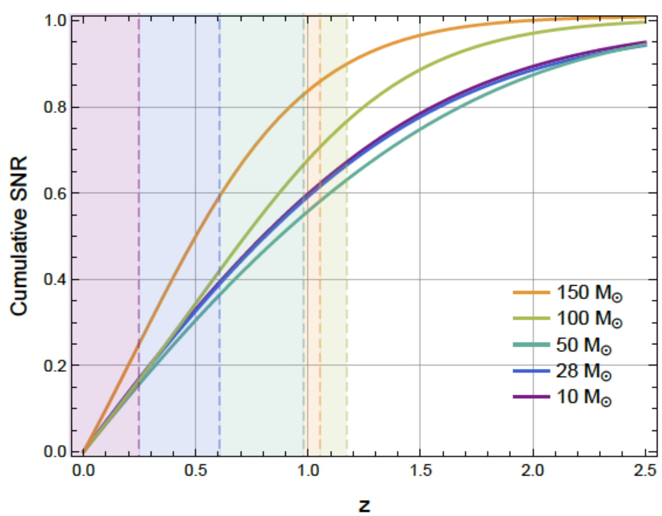
http://www.phys.lsu.edu/~rih/binsim/gallery.html

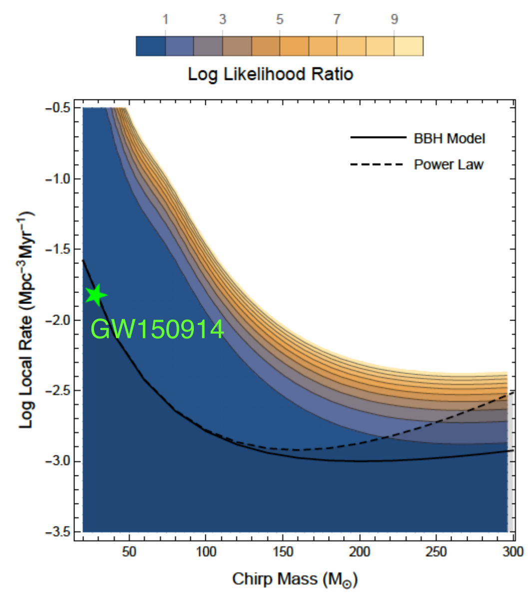
Contribution to a stochastic astrophysical background

- In addition to individual *foreground* events, we expect a *stochastic background* of many unresolved, distant events from all directions at essentially all times ("popcorn noise").
- There will be a (redshifted) cutoff frequency, depending on the average chirp mass of the systems that dominate this background.
- For low mass systems, foreground events account for only a small fraction of the total SNR in the stochastic signal.
- The background associated with events like GW150914 may be marginally detectable

- Caltech LIGO: Tom Callister

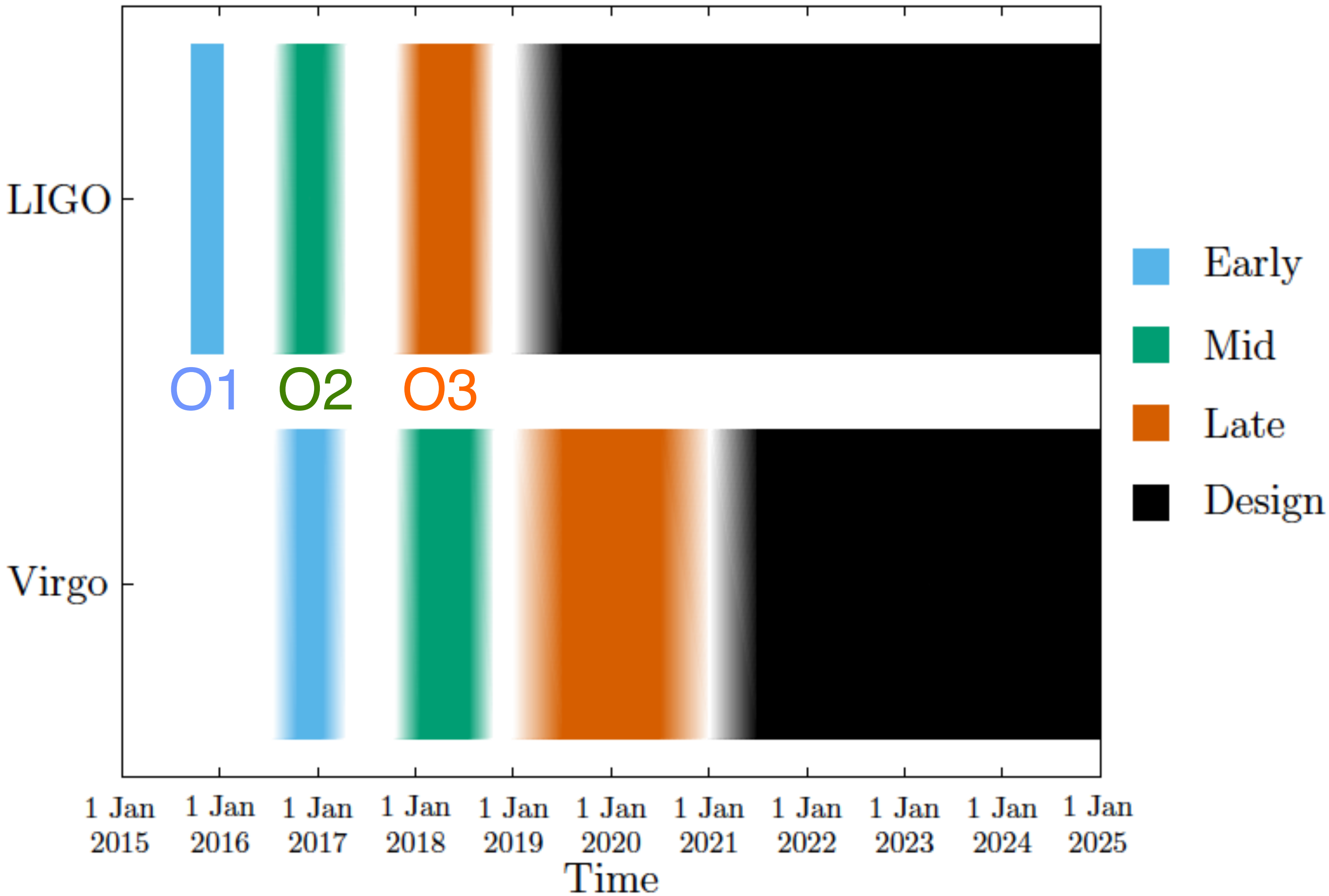






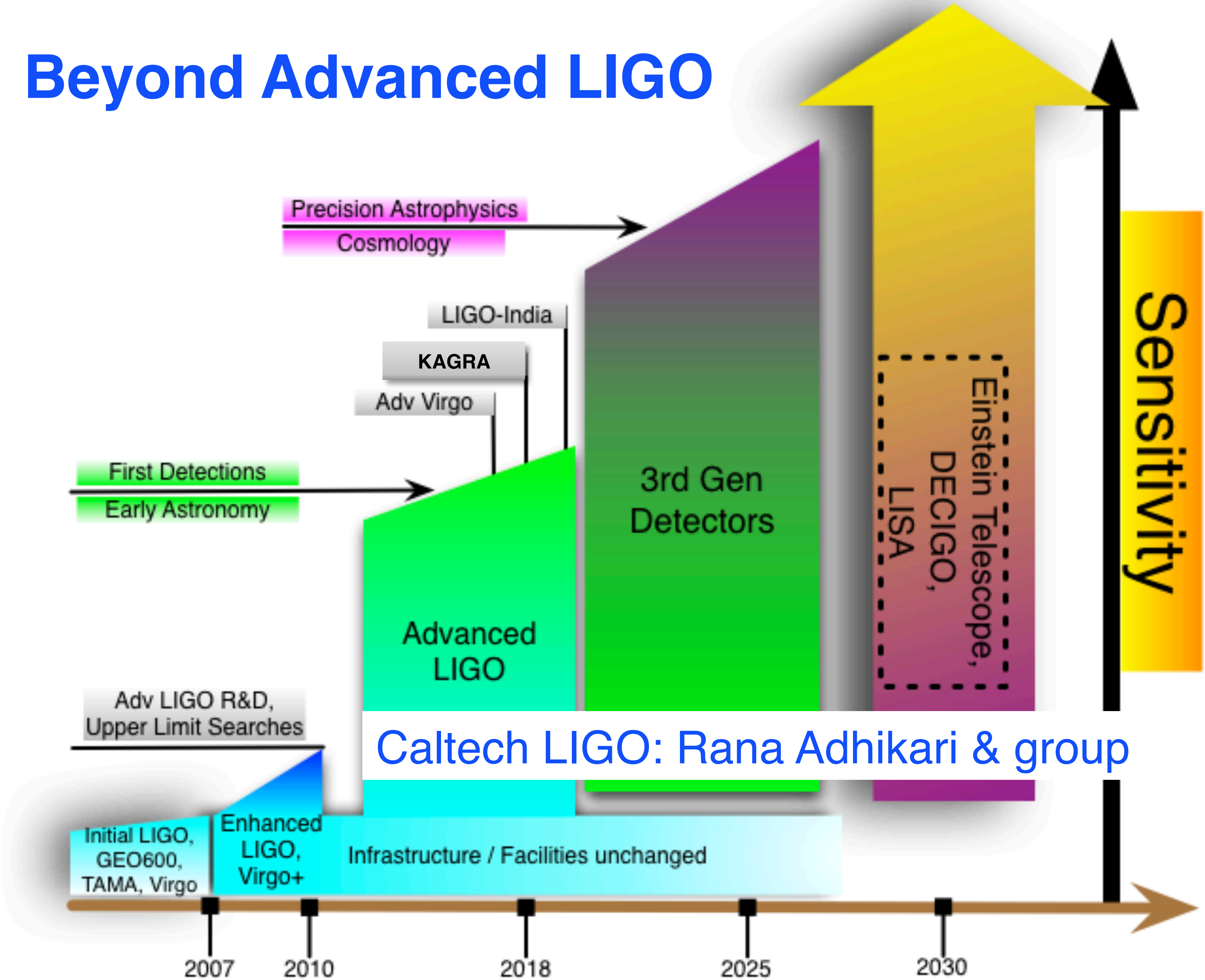
https://dcc.ligo.org/LIGO-P1500222/public/main ; Tom Callister, in prep

near-term future – very preliminary plan



http://relativity.livingreviews.org/Articles/lrr-2016-1/

Beyond Advanced LIGO



Caltech LIGO: Rana Adhikari & group

More info: papers.ligo.org (much more to come!)

LIGO
LASER INTERFEROMETER GRAVITATIONAL-WAVE OBSERVATORY

Observation of Gravitational Waves from a Binary Black Hole Merger

Download Paper

Abstract:

On September 14, 2015 at 09:50:45 UTC the two detectors of the Laser Interferometer Gravitational-Wave Observatory simultaneously observed a transient gravitational-wave signal. The signal sweeps upwards in frequency from 35 to 250 Hz with a peak gravitational-wave strain of 1.0×10^{-21}. It matches the waveform predicted by general relativity for the inspiral and merger of a pair of black holes and the ringdown of the resulting single black hole. The signal was observed with a matched-filter signal-to-noise ratio of 24 and a false alarm rate estimated to be less than 1 event per 203 000 years, equivalent to a significance greater than 5.1σ. The source lies at a luminosity distance of 410^{+160}_{-180} M_{pc} corresponding to a redshift $z = 0.09^{+0.03}_{-0.04}$. In the source frame, the initial black hole masses are $36^{+5}_{-4} M_\odot$ and $29^{+4}_{-4} M_\odot$, and the final black hole mass is $62^{+4}_{-4} M_\odot$, with $3.0^{+0.5}_{-0.5} M_\odot c^2$ radiated in gravitational waves. All uncertainties define 90% credible intervals. These observations demonstrate the existence of binary stellar-mass black hole systems. This is the first direct detection of gravitational waves and the first observation of a binary black hole merger.

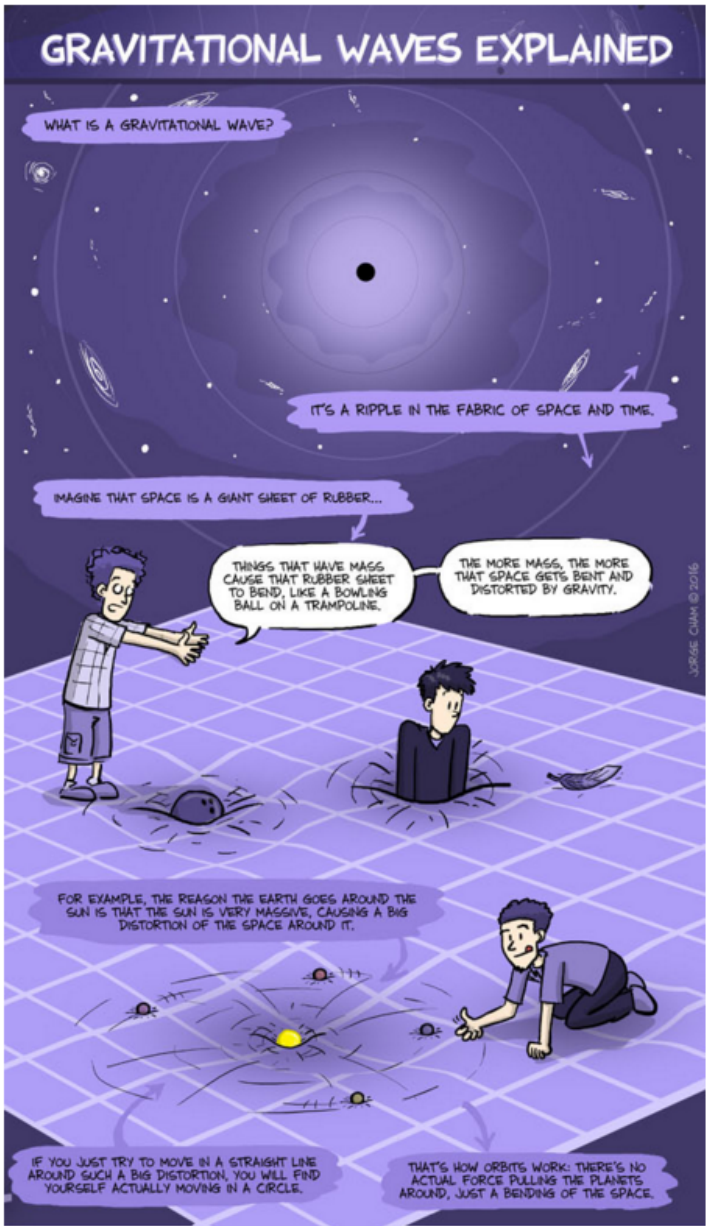
Related Papers:

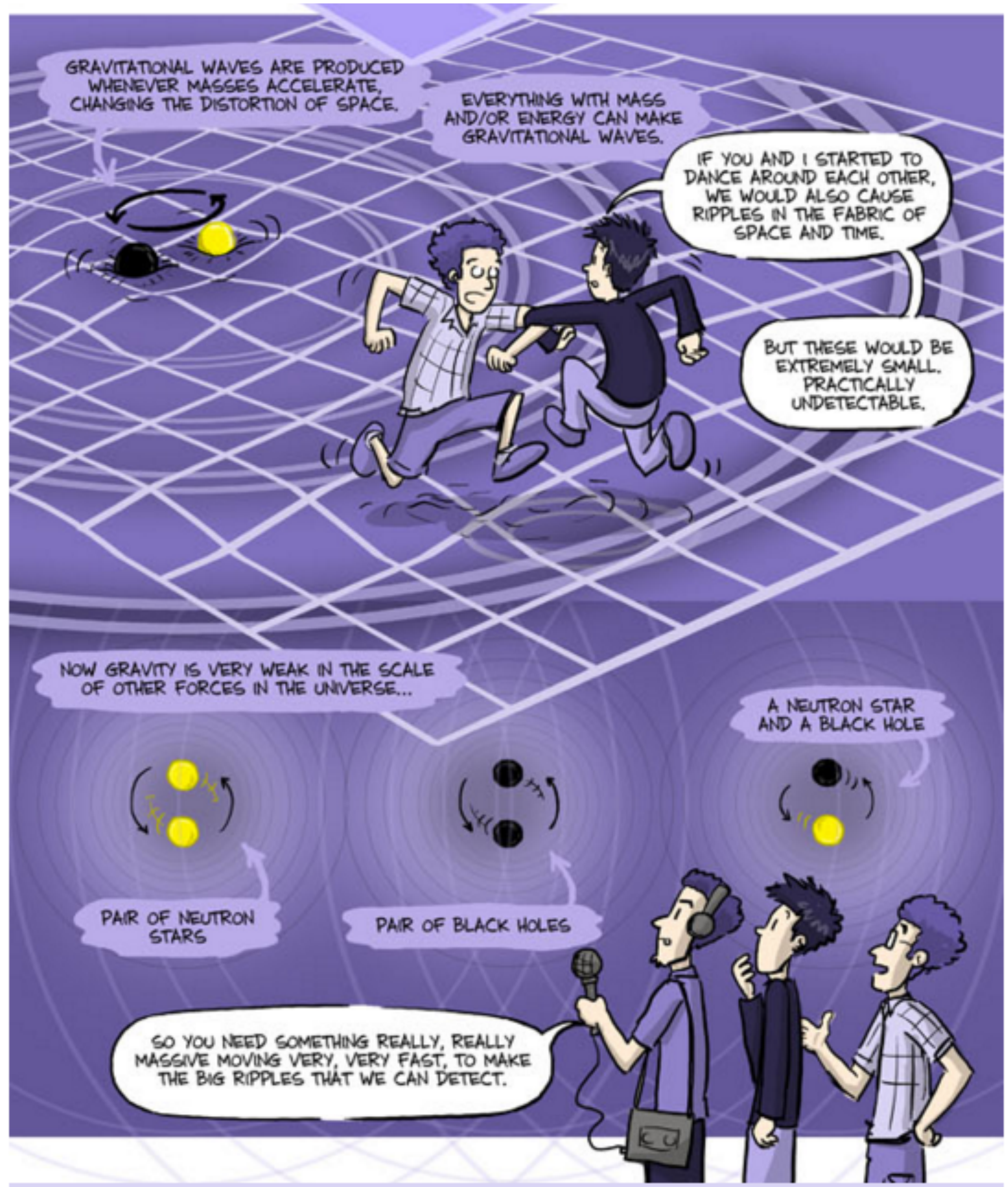
- LIGO-P1500229: Observing gravitational-wave transient GW150914 with minimal assumptions
- LIGO-P1500269: GW150914: First results from the search for binary black hole coalescence with Advanced LIGO
- LIGO-P1500218: Properties of the binary black hole merger GW150914
- LIGO-P1500217: The Rate of Binary Black Hole Mergers Inferred from Advanced LIGO Observations Surrounding GW150914
- LIGO-P1500262: Astrophysical Implications of the Binary Black-Hole Merger GW150914
- LIGO-P1500213: Tests of general relativity with GW150914
- LIGO-P1500222: GW150914: Implications for the stochastic gravitational-wave background from binary black holes
- LIGO-P1500248: Calibration of the Advanced LIGO detectors for the discovery of the binary black-hole merger GW150914
- LIGO-P1500238: Characterization of transient noise in Advanced LIGO relevant to gravitational wave signal GW150914
- LIGO-P1500227: Localization and broadband follow-up of the gravitational-wave transient GW150914
- LIGO-P1500271: High-energy Neutrino follow-up search of Gravitational Wave Event GW150914 with IceCube and ANTARES
- LIGO-P1500237: GW150914: The Advanced LIGO Detectors in the Era of First Discoveries

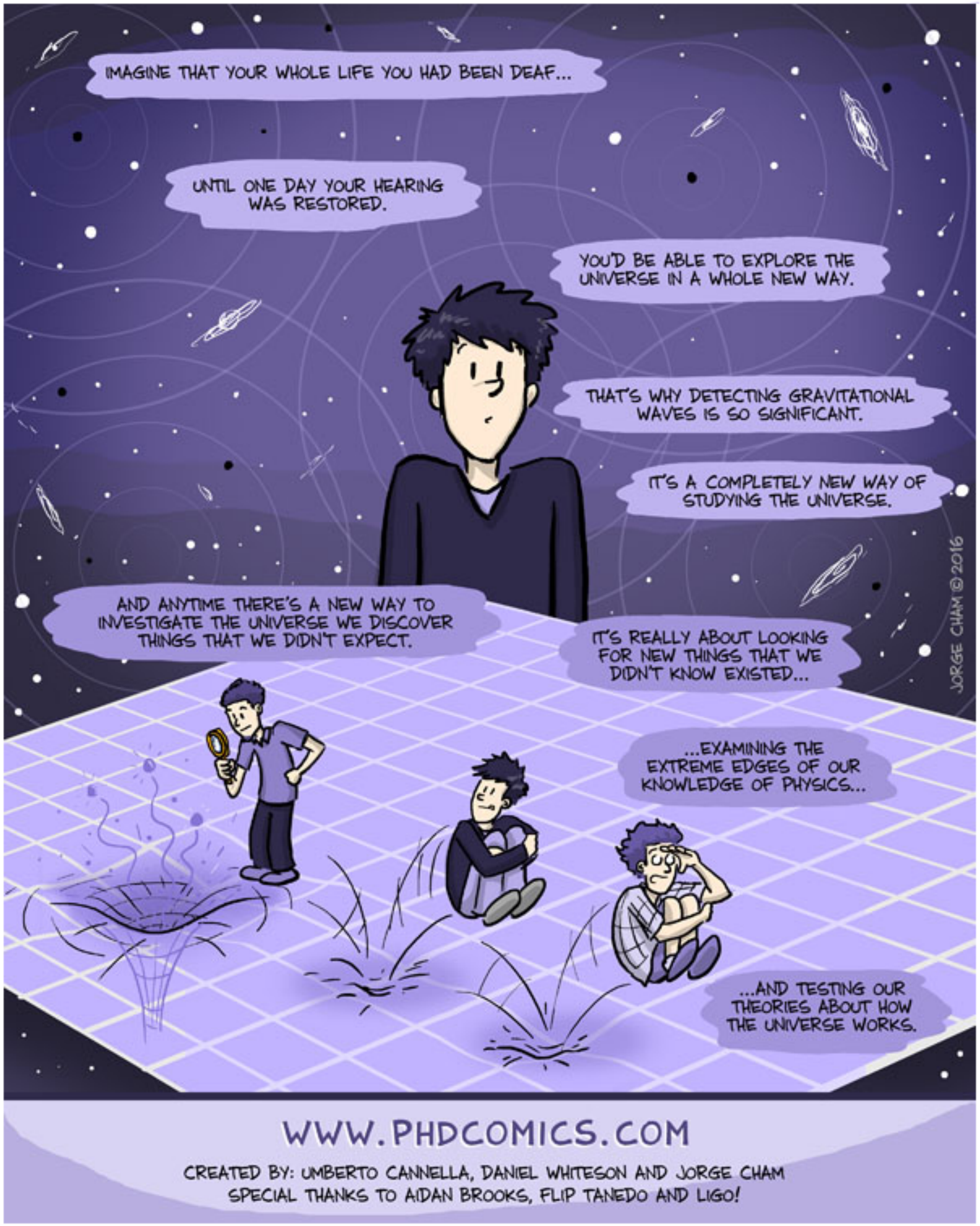
LIGO

If none of this was comprehensible… see http://phdcomics.com/comics.php?f=1853 (over 1 million page views!)

LSC






IMAGINE THAT YOUR WHOLE LIFE YOU HAD BEEN DEAF...
UNTIL ONE DAY YOUR HEARING WAS RESTORED.
YOU'D BE ABLE TO EXPLORE THE UNIVERSE IN A WHOLE NEW WAY.
THAT'S WHY DETECTING GRAVITATIONAL WAVES IS SO SIGNIFICANT.
IT'S A COMPLETELY NEW WAY OF STUDYING THE UNIVERSE.
AND ANYTIME THERE'S A NEW WAY TO INVESTIGATE THE UNIVERSE WE DISCOVER THINGS THAT WE DIDN'T EXPECT.
IT'S REALLY ABOUT LOOKING FOR NEW THINGS THAT WE DIDN'T KNOW EXISTED...
...EXAMINING THE EXTREME EDGES OF OUR KNOWLEDGE OF PHYSICS...
...AND TESTING OUR THEORIES ABOUT HOW THE UNIVERSE WORKS.
JORGE CHAM © 2016
WWW.PHDCOMICS.COM
CREATED BY: UMBERTO CANNELLA, DANIEL WHITESON AND JORGE CHAM
SPECIAL THANKS TO AIDAN BROOKS, FLIP TANEDO AND LIGO!

Bumper stickers

Gravitational Wave Detection!

Gravitational Wave Detection
is HARD!

PRECISION Gravitational Wave
science is HARDER!

Physics and astrophysics with gravitational waves

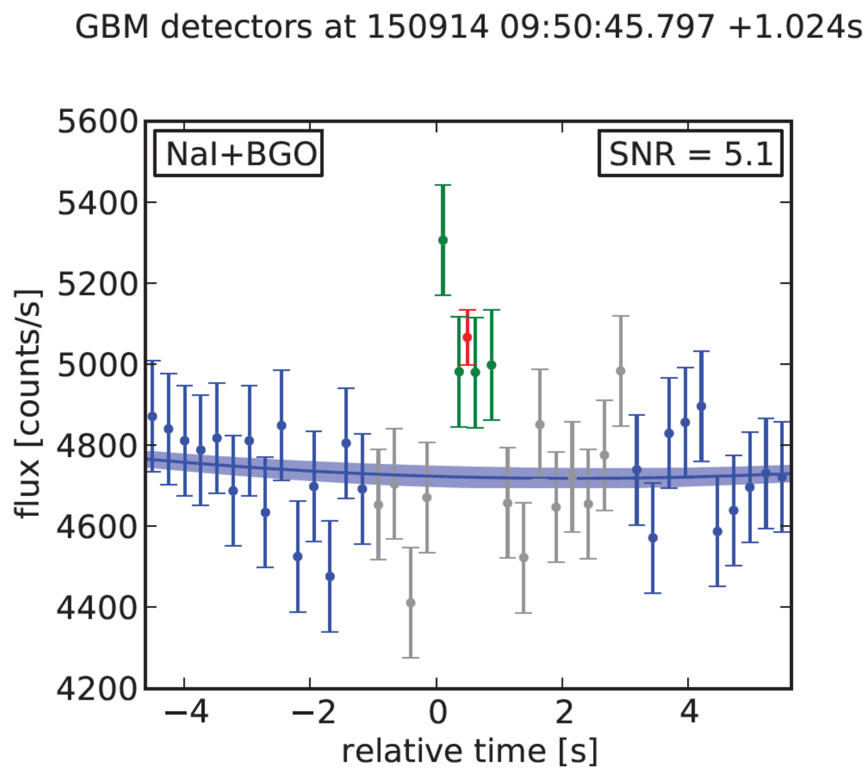
The advanced GW detector era has begun!

- **The exploration of the GW sky;**
- **unique tests of General Relativity in the strong-field, highly non-linear and dynamical regime;**
- **joint observations and discoveries with EM and neutrino telescopes;**
- **and a rich new branch of astrophysics.**

But most of all, we look forward to …

the unexpected!

Fermi GBM around the time of the event



FAP=0.0022

Fig. 2.— Count rates detected as a function of time relative to the start of GW150914-GBM, ∼0.4 s after the GW event GW150914, weighted and summed to maximize signal-to-noise for a modeled source. CTIME time bins are 0.256 s wide. The blue data points are used in the background fit. The green points are the counts in the time period determined to be significant, the grey points are outside this time period, and the red points show the 1.024 s average over the green points. For a